



The Green Bank Ammonia Survey: First Results of NH₃ Mapping of the Gould Belt

Citation

Friesen, Rachel K., Jaime E. Pineda, Erik Rosolowsky, Felipe Alves, Ana Chacón-Tanarro, Hope How-Huan Chen, Michael Chun-Yuan Chen, et al. 2017. "The Green Bank Ammonia Survey: First Results of NH₃ Mapping of the Gould Belt." *The Astrophysical Journal* 843 (1) (July 3): 63. doi:10.3847/1538-4357/aa6d58.

Published Version

10.3847/1538-4357/aa6d58

Permanent link

<http://nrs.harvard.edu/urn-3:HUL.InstRepos:37308771>

Terms of Use

This article was downloaded from Harvard University's DASH repository, and is made available under the terms and conditions applicable to Open Access Policy Articles, as set forth at <http://nrs.harvard.edu/urn-3:HUL.InstRepos:dash.current.terms-of-use#OAP>

Share Your Story

The Harvard community has made this article openly available.
Please share how this access benefits you. [Submit a story](#).

[Accessibility](#)

THE GREEN BANK AMMONIA SURVEY (GAS): FIRST RESULTS OF NH₃ MAPPING THE GOULD BELT

RACHEL K. FRIESEN¹ AND JAIME E. PINEDA²
(CO-PIs)

ERIK ROSOLOWSKY,³ FELIPE ALVES,² ANA CHACÓN-TANARRO,² HOPE HOW-HUAN CHEN,⁴ MICHAEL CHUN-YUAN CHEN,⁵
JAMES DI FRANCESCO,^{5,6} JARED KEOWN,⁵ HELEN KIRK,⁶ ANNA PUNANOVA,² YOUNGMIN SEO,⁷ YANCY SHIRLEY,⁸
ADAM GINSBURG,⁹ CHRISTINE HALL,¹⁰ STELLA S. R. OFFNER,¹¹ AYUSHI SINGH,¹² HÉCTOR G. ARCE,¹³ PAOLA CASELLI,²
ALYSSA A. GOODMAN,⁴ PETER G. MARTIN,¹⁴ CHRISTOPHER MATZNER,¹² PHILIP C. MYERS,⁴ AND ELENA REDAELLI^{2,15}
(THE GAS COLLABORATION)

¹*Dunlap Institute for Astronomy & Astrophysics, University of Toronto, 50 St. George Street, Toronto, Ontario, Canada M5S 3H4*

²*Max-Planck-Institut für extraterrestrische Physik, Giessenbachstrasse 1, 85748 Garching, Germany*

³*Department of Physics, University of Alberta, Edmonton, AB, Canada*

⁴*Harvard-Smithsonian Center for Astrophysics, 60 Garden St., Cambridge, MA 02138, USA*

⁵*Department of Physics and Astronomy, University of Victoria, 3800 Finnerty Road, Victoria, BC, Canada V8P 5C2*

⁶*Herzberg Astronomy and Astrophysics, National Research Council of Canada, 5071 West Saanich Road, Victoria, BC, V9E 2E7, Canada*

⁷*Jet Propulsion Laboratory, NASA, 4800 Oak Grove Dr, Pasadena, CA 91109, USA*

⁸*Steward Observatory, 933 North Cherry Avenue, Tucson, AZ 85721, USA*

⁹*National Radio Astronomy Observatory, Socorro, NM 87801, USA*

¹⁰*Department of Physics, Engineering Physics & Astronomy, Queen's University, Kingston, Ontario, Canada K7L 3N6*

¹¹*Department of Astronomy, University of Massachusetts, Amherst, MA 01003, USA*

¹²*Department of Astronomy & Astrophysics, University of Toronto, 50 St. George Street, Toronto, Ontario, Canada M5S 3H4*

¹³*Department of Astronomy, Yale University, P.O. Box 208101, New Haven, CT 06520-8101, USA*

¹⁴*Canadian Institute for Theoretical Astrophysics, University of Toronto, 60 St. George St., Toronto, Ontario, Canada, M5S 3H8*

¹⁵*Dipartimento di Fisica & Astronomia, Università degli Studi di Bologna, Viale Berti Pichat, 6/2, I - 40127 Bologna, Italy*

(Accepted March 28, 2017)

Submitted to ApJS

Abstract

We present an overview of the first data release (DR1) and first-look science from the Green Bank Ammonia Survey (GAS). GAS is a Large Program at the Green Bank Telescope to map all Gould Belt star-forming regions with $A_V \gtrsim 7$ mag visible from the northern hemisphere in emission from NH₃ and other key molecular tracers. This first release includes the data for four regions in Gould Belt clouds: B18 in Taurus, NGC 1333 in Perseus, L1688 in Ophiuchus, and Orion A North in Orion. We compare the NH₃ emission to dust continuum emission from *Herschel*, and find that the two tracers correspond closely. NH₃ is present in over 60 % of lines-of-sight with $A_V \gtrsim 7$ mag in three of the four DR1 regions, in agreement with expectations from previous observations. The sole exception is B18, where NH₃ is detected toward ~ 40 % of lines-of-sight with $A_V \gtrsim 7$ mag. Moreover, we find that the NH₃ emission is generally extended beyond the typical 0.1 pc length scales of dense cores. We produce maps of the gas kinematics, temperature, and NH₃ column densities through forward modeling of the hyperfine structure of the NH₃ (1,1) and (2,2) lines. We show that the NH₃ velocity dispersion, σ_v , and gas kinetic temperature, T_K , vary systematically between the regions included in this release, with an increase in both the mean value and spread of σ_v and T_K with increasing star formation activity. The data presented in this paper are publicly available.

Keywords: stars:formation — ISM:molecules — ISM:individual (OrionA molecular complex) — ISM:individual (NGC1333) — ISM:individual (B18) — ISM:individual (L1668)

1. INTRODUCTION

The past several years have seen tremendous advancements in our ability to characterize the structure of nearby molecular clouds and the substructures in which dense star-forming cores are born. Within 500 pc of the Sun, nearly all the ongoing, predominantly low-mass star formation is contained within a ring of young stars and star-forming regions named the Gould Belt. These nearby star-forming clouds range in size and activity, from the nearly inactive Pipe Nebula, through the star-forming but largely quiescent Taurus molecular cloud, to stellar group and cluster-forming clouds like Perseus and Ophiuchus. The Orion molecular cloud is our nearest example of a high-mass star-forming region. This variation in star formation activity, coupled with their relatively nearby distances, make the Gould Belt clouds excellent survey targets to understand how star formation proceeds in different environments, and to constrain star formation theories and simulations through detailed observations.

Consequently, the Gould Belt star-forming regions have been surveyed by several Key Programs in the infrared and sub-millimeter continuum. The infrared surveys have characterized the embedded young stellar populations, identifying the loci of active star formation and constraining the lifetimes of the starless and embedded stages of star-forming cores (the *Spitzer* Gould Belt, *Spitzer* Orion and c2d surveys; Enoch et al. 2008; Harvey et al. 2008; Kirk et al. 2009, Megeath et al. 2012, and Evans et al. 2003, 2009; Dunham et al. 2014). At longer wavelengths, the emission from cold dust within dense, star-forming cores, and the larger filaments and clumps that form the cores and young stellar objects (YSOs), are being surveyed extensively through Gould Belt Legacy surveys with the *Herschel Space Observatory* (*Herschel* GBS; André et al. 2010) and the James Clerk Maxwell Telescope (JCMT GBS; Ward-Thompson et al. 2007). The *Herschel* GBS has revealed the striking prevalence of filaments within star-forming regions as seen in the dust emission (André et al. 2010; Könyves et al. 2010). While the presence of filaments in star-forming regions was not a new discovery (e.g., Schneider & Elmegreen 1979), their ubiquity at both low and high column densities suggests that they are integral in the build-up of the required mass to form stars. Furthermore, the *Herschel* GBS has found gravitationally unstable cores predominantly within (apparently) gravitationally unstable filaments. Previous studies have shown that dense molecular cores are mostly found above a visual extinction threshold of $A_V > 5 \sim 7$ mag ($N(\text{H}_2) \sim 5 - 7 \times 10^{21} \text{ cm}^{-2}$; e.g., Onishi et al. 1998; Johnstone et al. 2004; Kirk et al. 2006). With a suggested characteristic filament width of ~ 0.1 pc (Arzoumanian et al. 2011, but see Panopoulou et al. 2017), this extinction threshold agrees within a factor of two with the critical mass per unit length of an isothermal,

thermally-supported cylinder ($M_{\text{line,crit}} \sim 16 M_\odot \text{ pc}^{-1}$ at 10 K; Inutsuka & Miyama 1997; Hocuk et al. 2016). Therefore, understanding filament instability might be central to understanding the formation of dense cores and, consequently, stars.

If the massive and long filaments found by the *Herschel* GBS are an important piece of the star formation puzzle (André et al. 2014), then we must understand what are their main formation mechanisms, evolutionary drivers, and fragmentation triggers. The power of the legacy surveys described above lies in the large areal coverage and consistency in observing strategies between nearby molecular clouds. A major gap in the present data is a comparable survey to characterize the dense gas properties. Continuum data show the dust column density structure, but do not provide information about the gas kinematics. Kinematics and gas temperatures are key to understanding the history and future fate of star-forming material.

Kinematic studies of filaments are relatively few (e.g., Pineda et al. 2011; Hacar & Tafalla 2011; Hacar et al. 2013; Pineda et al. 2015; Hacar et al. 2016; Henshaw et al. 2016a). The CARMA Large Area Star Formation Survey (Storm et al. 2014) targeted several regions within nearby Gould Belt clouds, at higher resolution but on a smaller scale relative to the JCMT and *Herschel* continuum surveys. The JCMT GBS has produced the widest survey thus far, using observations of CO (3-2) and its isotopologues to trace moderately dense gas (10^3 cm^{-3} , e.g., Graves et al. 2010). CO, however, is frequently very optically thick and the gas phase abundance can be severely affected by depletion onto dust grains (e.g., Caselli et al. 1999; Tafalla et al. 2002; Christie et al. 2012) in star-forming filaments and cores, and therefore it is not a useful tracer of the dense gas.

The lower NH_3 metastable inversion transitions primarily trace gas of density $n \geq 2 \times 10^3 \text{ cm}^{-3}$ (for gas at 10 K, Shirley 2015). Furthermore, NH_3 does not typically suffer from depletion. Since NH_3 observations allow us to identify actual high volume density features rather than column density peaks, NH_3 is then an ideal tracer of the hierarchical nature of star-forming regions, and a powerful probe of the kinematics of star-forming filaments and cores.

Past NH_3 observations of dense gas were primarily pencil-beam pointings or small maps around interesting targets (e.g., Jijina et al. 1999; Rosolowsky et al. 2008). The NH_3 -identified cores are generally cold and quiescent (Goodman et al. 1998), either in early stages of collapse or near-critical equilibrium (Benson & Myers 1983), and often show significant velocity gradients (Goodman et al. 1993). The local environment influences key core properties, such as the ratio of thermal vs. non-thermal motions and temperature (Myers et al. 1991; Jijina et al. 1999; Foster et al. 2009), both important parameters for investigations of core stability

Table 1. GAS DR1 region data

Cloud	Region	Distance (pc)	Number of footprints ^a
Taurus	B18	135 ± 20^b	11
Ophiuchus	L1688	137.3 ± 6^c	10
Perseus	NGC 1333	260 ± 26^b	8
Orion	Orion A-North	414 ± 7^d	12

^aEach footprint is $10' \times 10'$.

^bSchlaflly et al. (2014)

^cOrtiz-León et al. (2017)

^dMenten et al. (2007)

Table 2. GAS VEGAS spectral line setup

Species	Transition	Rest frequency (MHz)	Reference
NH ₃	(1, 1)	23694.4955	
NH ₃	(2, 2)	23722.6336	
NH ₃	(3, 3)	23870.1296	
HC ₅ N	9 – 8	23963.9010	LOVAS
HC ₇ N	21 – 20	23687.8974(6)	CDMS
HC ₇ N	22 – 21	24815.8772(6)	CDMS
C ₂ S ^a	2 ₁ – 1 ₀	22344.030(1)	CDMS

^aObserved in central KPFA beam only

and collapse. Larger, more sensitive maps are needed to relate the physical properties of the cores to their surrounding environment. These recent studies have highlighted the sharp transition between turbulent and thermal line widths in dense cores (Pineda et al. 2010; Seo et al. 2015), identified incongruities between structures identified via continuum and line emission (Friesen et al. 2009), shown that cores in more active environments, such as Orion, tend to be more unstable, despite having similar internal motions (Li et al. 2013), identified large-scale gravitational instability in a young filamentary cluster (Friesen et al. 2016), and revealed evidence for filamentary accretion in Serpens South (Friesen et al. 2013, also in N₂H⁺ by Kirk et al. 2013).

Here, we describe the Green Bank Ammonia Survey (GAS, co-PIs: R. Friesen and J. E. Pineda), an am-

bitious Large Project to survey all the major, nearby ($120 < d < 500$ pc), northern Gould Belt star-forming regions with visual extinctions $A_V \gtrsim 7$ (matching the extinction threshold for dense cores in continuum studies e.g., Johnstone et al. 2004; Enoch et al. 2007; Lada et al. 2010) in emission from NH₃, as well as the carbon-chain molecules C₂S, HC₅N, and HC₇N. Within this distance range, the GBT beam at 23 GHz (32'' FWHM) subtends 0.02 pc – 0.08 pc, able to resolve both the Jeans length (0.12 pc for $n = 10^4$ cm⁻³ and $T = 10$ K) and the ~ 0.1 pc typical filament width suggested by Arzoumanian et al. (2011), and is furthermore well-matched to the resolution of *Herschel* SPIRE at 500 μ m ($\theta_{\text{beam}} = 36''$ FWHM; Griffin et al. 2010). This coordinated, large-scale survey of NH₃ in the Gould Belt clouds will play a key role in understanding the evolu-

tion of dense gas in star-forming regions as a function of environment, expanding in a consistent way the NH_3 studies described above, and will complement and enhance the Gould Belt continuum surveys.

In this paper, we present the first data release (hereafter referred to as DR1) of a subset of the survey targets, describing in detail the calibration, imaging, and analysis pipelines, and discussing some general trends in the data. The target regions presented are Barnard 18 in Taurus (hereafter B18), NGC 1333 in Perseus, L1688 in Ophiuchus, and Orion A (North). Future research will focus on the key scientific goals of GAS, including analyzing the structure and stability of the dense gas, the dissipation of turbulence from large to small scales, and the evolution of angular momentum in dense cores.

In § 2, we discuss the source selection, observations, data calibration and imaging. In § 3, we describe in detail the hyperfine modeling of the NH_3 emission and the determination of key physical parameters of the emitting gas. In § 4, we present the moment maps and noise properties of the regions included in this first data release, as well as the results from the NH_3 line modeling. We present a summary in § 5.

2. DATA

2.1. Source selection and description

The goal of GAS was to map NH_3 emission toward all areas within the nearby ($d < 500$ pc), GBT-visible, star-forming molecular clouds of the Gould Belt with extinctions $A_V > 7$ mag [$N(\text{H}_2) \gtrsim 6.7 \times 10^{21} \text{ cm}^{-2}$, assuming $N(\text{H}_2) = 9.4 \times 10^{20} \text{ cm}^{-2} (A_V \text{ mag}^{-1})$] (Johnstone et al. 2004; Enoch et al. 2007; Heiderman et al. 2010; Lada et al. 2010). The initial map extents were identified in each cloud using either publicly released *Herschel* GBS 500 μm continuum data, unreleased JCMT GBLs 850 μm continuum data, or extinction maps derived from 2MASS data (Dobashi et al. 2005; Ridge et al. 2006). Further refinement of mapping targets was done as new data were publicly released (e.g., *Herschel* GBS).

For our initial observing season we focused on four regions that span a range of physical environments and star formation activity: B18, NGC 1333, L1688, and Orion A (North). These regions comprise the first data release (DR1) of GAS.

B18 is a filamentary structure located in the southeast of the Taurus molecular cloud, and runs roughly parallel in projection with the molecular gas structure containing Heiles Cloud 2 and the B213/L1495 filament. B18 has a projected length of $\sim 3^\circ$ (~ 7 pc) with a total mass of $380 - 440 M_\odot$ (Heyer et al. 1987; Mizuno et al. 1995), giving a mean gas number density $n \sim 300 \text{ cm}^{-3}$, and contains ~ 9 molecular clumps as traced by CO (Heyer 1988). Based on the numbers of associated young stars, the star formation efficiency (SFE), where $SFE = \text{current } M_*/M_{\text{gas}}$, is $\sim 1 - 2 \%$ (Mizuno et al. 1995).

NGC 1333 is the most active star-forming region within the Perseus molecular cloud and contains approximately $450 M_\odot$ of molecular gas (Warin et al. 1996) and ~ 150 young stars (Gutermuth et al. 2008) in a region $\sim 1 \text{ pc}^2$. The mean gas density is therefore a few $\times 10^3 \text{ cm}^{-3}$. Rough estimates suggest the star formation efficiency is $\sim 13 - 15 \%$ in the region over the past 10^6 years (Walsh et al. 2006; Jørgensen et al. 2008).

L1688 is the centrally-concentrated, dense hub of the Ophiuchus molecular cloud and is $\sim 1 - 2$ pc in radius and $\sim 2 \times 10^3 M_\odot$ in mass (Loren 1989). Similarly to NGC 1333, the mean gas density is approximately a few $\times 10^3 \text{ cm}^{-3}$. L1688 contains regions of extremely high visual extinction, with $A_V \sim 50 - 100$ mag (e.g., Wilking & Lada 1983), and over 300 young stellar objects (Wilking et al. 2008). The star formation efficiency of the dense gas cores, identified in submillimeter continuum emission from dust, is $\sim 14 \%$ (Jørgensen et al. 2008).

Orion A (North), an integral-shaped filamentary structure of dense gas that extends both north and south of the Orion Nebula and its associated young stellar cluster, including the Trapezium stars, is our nearest example of ongoing high-mass star formation. The filament is characterized by varying physical conditions along the ~ 12 pc length mapped for DR1. A compact, narrow ridge north of the Orion Nebula, the filament becomes wider and less dense toward the south, and contains a total mass of $\sim 5 \times 10^3 M_\odot$ in the area mapped here (e.g., Bally et al. 1987; Bally 2008). Overall, the star formation efficiency in Orion A is $\sim 3 - 5 \%$ in gas with extinction $A_K > 2$, with significantly higher values (up to several tens of percent) found toward young protostellar groups and clusters (Megeath et al. 2015).

Table 1 lists the assumed distances and number of footprints (areal coverage of $10' \times 10'$) completed for the GAS DR1 regions.

2.2. Observations

For all regions, observations were performed using the seven-beam K-Band Focal Plane Array (KFPA) at the GBT, with the VErSatile GBT Astronomical Spectrometer (VEGAS) backend. GAS uses VEGAS configuration Mode 20, allowing eight separate spectral windows per KFPA beam, each with a bandwidth of 23.44 MHz and 4096 spectral channels, for a spectral resolution of 5.7 kHz, or $\sim 0.07 \text{ km s}^{-1}$ at 23.7 GHz. All spectral lines observed are listed in Table 2. Due to limitations on the maximum separation in GHz between spectral lines in a single VEGAS bank, the GAS setup includes six spectral lines observed in each KFPA beam, plus the $\text{C}_2\text{S } 2_1 - 1_0$ line in a single central beam. Observations were performed using in-band frequency switching to maximize on-source time, with a frequency throw of 4.11 MHz.

GAS maps were generally observed in $10' \times 10'$ footprints, scanning in Right Ascension (R.A.), with scan rows separated by $13''$ in Declination (Dec.) to en-

Table 3. GAS relative beam calibration

Pol	Beams						
	1	2	3	4	5	6	7
LL	0.894 (0.071)	0.865 (0.101)	0.847 (0.049)	0.772 (0.074)	0.925 (0.062)	0.677 (0.081)	0.851 (0.080)
RR	0.845 (0.060)	0.846 (0.072)	0.816 (0.074)	0.745 (0.077)	0.764 (0.057)	0.511 (0.046)	0.970 (0.081)

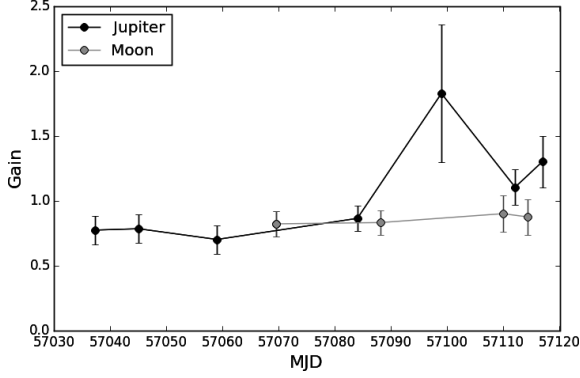


Figure 1. Average gains over all beams and polarizations for the NH_3 (1,1) spectral window as a function of date derived from observations of Jupiter and the Moon. Error bars show the $1 - \sigma$ variation. As discussed further in the text, the Jupiter observations after 57090 MJD show large offsets from previous Jupiter measurements and concurrent Moon measurements due to high winds on those dates.

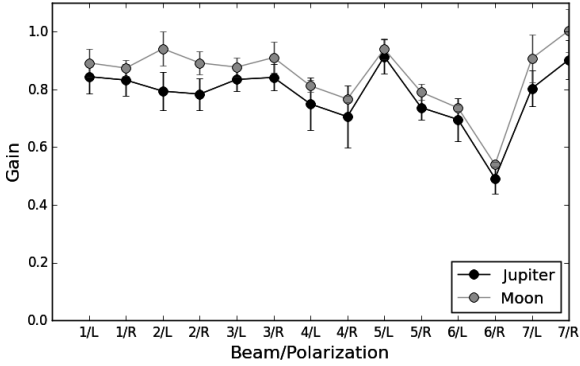


Figure 2. Beam gains for the NH_3 (1,1) spectral window averaged over all observing dates as a function of beam and polarization for Jupiter and the Moon. The observations of Jupiter during windy periods have been omitted. Error bars show the $1 - \sigma$ variation in the values. Although the gains vary between beams and polarizations, they remained consistent over the observing period.

sure Nyquist sampling. Each $10'$ square submap was observed once. To best match the observed regions to known sub-millimetre continuum emission, some regions included submaps of $10' \times 5'$, or $5' \times 10'$ (where scans were in Dec., spaced by $13''$ in R.A. to minimize overheads). For all maps, the telescope scan rate was $6''.2 \text{ s}^{-1}$, with data dumped every 1.044 s. A fast frequency-switching rate of 0.348 s was used to meet the proposed sensitivity per map in the time allotted, but resulted in a small fraction (approximately a few percent) of the data being blanked due to synthesizer settling.

2.3. Flux Calibration

Flux calibration was performed using observations of the Moon and Jupiter. Because the Moon is large relative to the KFPA footprint, simple on-target and off-target observations allowed relative beam calibrations to be performed. In angular size, Jupiter was comparable to the $32''$ GBT beam at 23.7 GHz. For these sources, nod observations were performed to alternate on- and off-source beams, cycling through all beams. Relative beam calibrations were then calculated in GBTIDL¹. Models for each solar system target, including the effects of solar illumination and apparent size, were used to determine the expected brightness temperature in the gain analysis².

The Moon was used as the primary calibrator when available, as the planetary calibrations could vary substantially when the weather was not ideal, particularly when the on-site winds were near the 5 m s^{-1} limits for KFPA operations at 23 GHz. This can be seen in Figure 1, where we show the beam- and polarization-averaged gains derived from observations of Jupiter and the Moon over the 15A observing semester. Jupiter and Moon calibrations agree well when the winds are low. The last three dates when Jupiter was observed had relatively high winds ($\gtrsim 5 \text{ m s}^{-1}$), lowering the pointing accuracy of the telescope, which is important to obtain an accurate flux calibration, but was ameliorated for the science observations by more frequent pointing calibration observations. Moreover, the actual telescope

¹ <http://gbtidl.nrao.edu/>

² <https://safe.nrao.edu/wiki/pub/GB/Knowledge/GBTMemos/GBTMemo273-11Feb25.pdf>

pointing is recorded in the position encoder positions, and these are the values used during the data cube gridding. Jupiter’s angular size over the observation dates was $\sim 40''$, only slightly larger than the $32''$ (FWHM) GBT beam. Small pointing errors introduced by high winds are thus able to impact greatly the derived gains. In contrast, we show in Figure 1 that the gains determined from the Moon were very stable over the 15A semester.

Figure 2 shows the gains as a function of beam and polarization, averaged over the observing semester. The final three Jupiter measurements that were impacted by winds have been omitted from the analysis. Although the gains vary between beams and polarizations, they remained consistent over the observing period with only small variation. The gains determined and used for calibration (next section) are listed in Table 3, along with the standard deviation of the measurements for each beam over the semester.

Two known targets with bright NH_3 (1,1) lines (L1489_PPC and L43) were also observed frequently in a single beam. We found that the peak line brightness varied by $< 10\%$ ($1-\sigma$) for these sources over the semester, in agreement with the calibration uncertainty estimated from our directly-measured beam gain variations.

2.4. Data reduction and calibration

All GAS data reduction and imaging codes (next section) are publicly available on GitHub³ and can be easily adapted for similar observations.

The data reduction was performed using a Python wrapper that called the GBT KFPA data reduction pipeline (Masters et al. 2011). The data were calibrated to main beam brightness temperature (T_{MB}) units using the relative gain factors for each of the beams and polarizations discussed above and listed in Table 3. Based on the 1σ variation in the beam gain calibrations, we estimate our calibration uncertainty to be $\sim 10\%$.

2.5. Imaging

We grid the On-The-Fly (OTF) observations following the prescription by Mangum et al. (2007). For each region, our mapping algorithm establishes a world coordinate system and corresponding pixel grid that covers the footprint of the observations, using a pixel size of $1/3$ of the beam FWHM (θ_{beam}) of the NH_3 (1,1) line. We then calculate the value of each map pixel as the weighted average of spectra at the center of the map pixel, using a truncated, tapered Bessel function for the weighting scheme. Each integration is given a weight, W , based on the location of the observation relative to the center of the pixel in the resulting map and the ob-

servations rms:

$$W(r) = \frac{J_1(\pi r/a)}{(\pi r/a)} e^{-r^2/b^2} \frac{1}{T_{\text{sys}}^2} \quad (1)$$

where $a = 1.55(\theta_{\text{beam}}/3)$, $b = 2.52(\theta_{\text{beam}}/3)$, r is the angular distance between the observed datum and the map datum, and T_{sys} is the recorded system temperature. Integrations with $r > \theta_{\text{beam}}$ are excluded from the average, as suggested by Mangum et al. (2007). For simplicity of comparison, we grid data for all lines to a common world coordinate system.

Online Doppler tracking is only applied for the NH_3 (1,1) line. For all other lines, we apply a frequency shift in the Fourier domain to each individual integration to align all spectra before convolution (gridding), i.e. Doppler tracking is done in software. Baseline subtraction is done in both the individual integrations (linear baseline) and in the individual pixels in the final data cubes (Legendre polynomials of order 3). In the latter case, sections of the spectrum with significant emission or within 2 km s^{-1} of the expected systemic velocity are excluded from the fit.

The KFPA on the GBT is a sparse array, with seven beams arranged in a hexagonal pattern, and beam centers separated by $94.88''$ on the sky. The mapping method described in § 2.2 ensured Nyquist-spaced sampling over $10' \times 10'$ footprints on the sky. At the edges of these footprints, coverage is variable, and the rms noise increases substantially as noted previously. For the final moment maps and parameters maps for each region, we mask pixels within 3 pixels of the map edges to remove the noisiest regions.

In Figure 3, we show histograms of the rms noise for all regions and lines observed. The rms noise map for the NH_3 (1,1) observations toward each DR1 region are presented in Appendix A.

In all regions, the noise histograms have strong peaks that agree well for the NH_3 (1,1), (2,2), (3,3), HC_5N 9 – 8, and HC_7N 21 – 10 observations, with typical 1σ noise values of $\lesssim 0.1\text{ K}$ per velocity channel across the DR1 regions. Note that the typical noise in individual map footprints can vary slightly depending on the conditions in which the region was mapped. Figure Set A show that in general this variation is minimal. The C_2S observations have greater rms noise values since C_2S was observed with a single beam rather than all seven beams in the KFPA. The HC_7N 22 – 21 data are noisier than the data of other lines observed with all KFPA beams, a systematic difference seen over all observations.

3. NH_3 LINE ANALYSIS

3.1. Line fitting

We calculate the properties of the column of ammonia molecules using a forward-modeling approach in the *pyspeckit* package (Ginsburg & Mirocha 2011), which builds from the works of Rosolowsky et al. (2008) and

³ <https://GitHub.com/GBTAmmoniaSurvey/GAS>

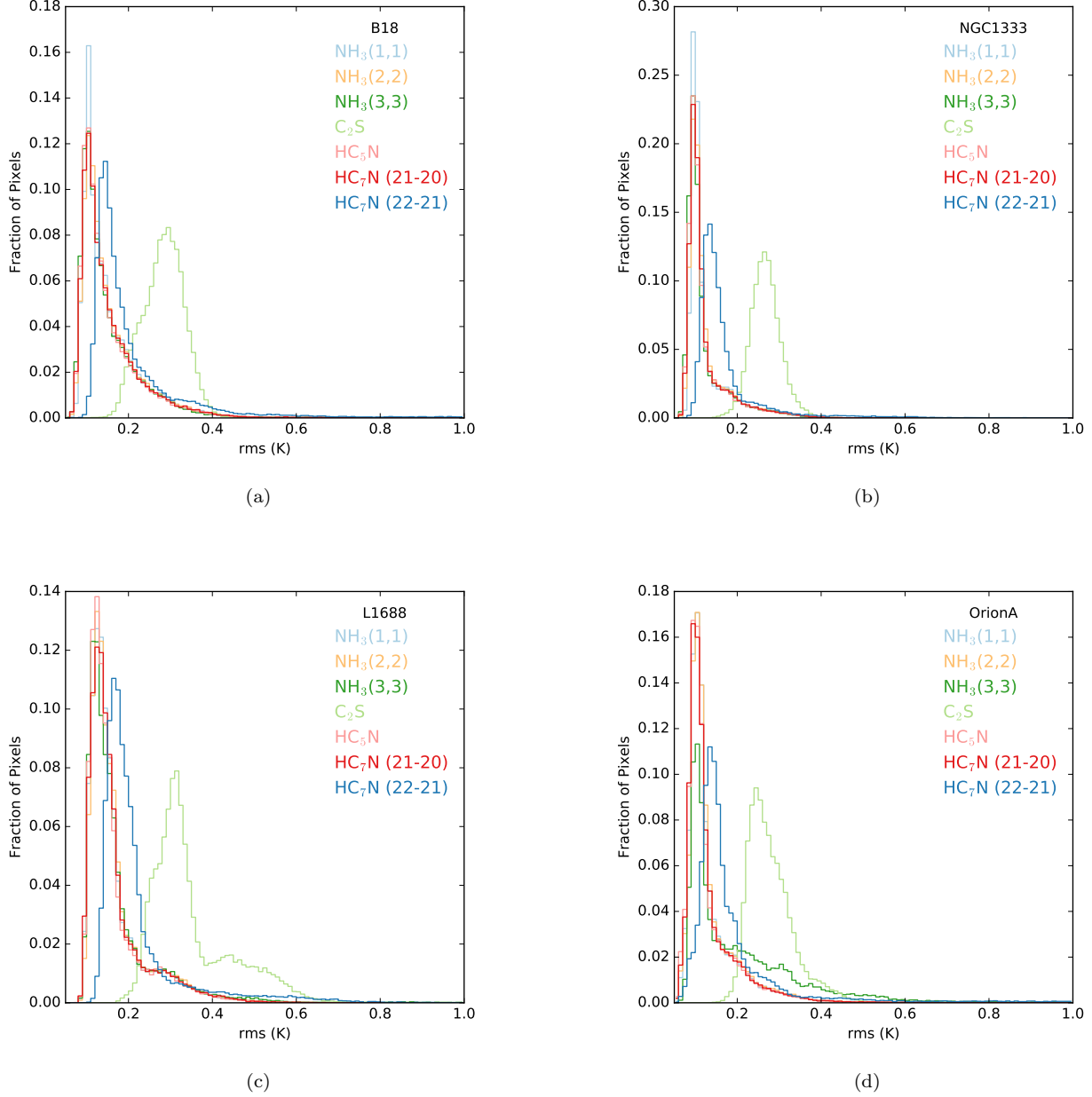


Figure 3. Histograms of the rms noise maps for all lines observed towards the DR1 regions, where each observed transition is shown in a different color. The rms is on the T_{MB} scale.

Friesen et al. (2009) relying on the theoretical framework laid out in Mangum & Shirley (2015). Under this approach, we consider a beam-filling slab of ammonia gas with a set of physical properties including para-ammonia column density ($N_{\text{p-NH}_3} \equiv N$ for notational compactness in this section), kinetic temperature (T_K), velocity dispersion σ_v , and line of sight velocity in the (kinematic) LSR frame (v_{LSR}). While we observe the $(J, K) = (3, 3)$ line of ortho-ammonia, we do not directly incorporate any information from this line into our fit

because of unknown ortho-to-para ratios that might differ from LTE (e.g., Faure et al. 2013).

Given a total column density for the molecule and a velocity dispersion for the slab, we use standard relations to determine the fraction of molecules found in each of the states, namely (J, K) rotational states of the species and the upper (u , antisymmetric inversion symmetry) and lower (l , symmetric inversion symmetry) within those states. Our model incorporates three distinct temperatures. The kinetic temperature T_K describes the velocity distributions of the particles in the

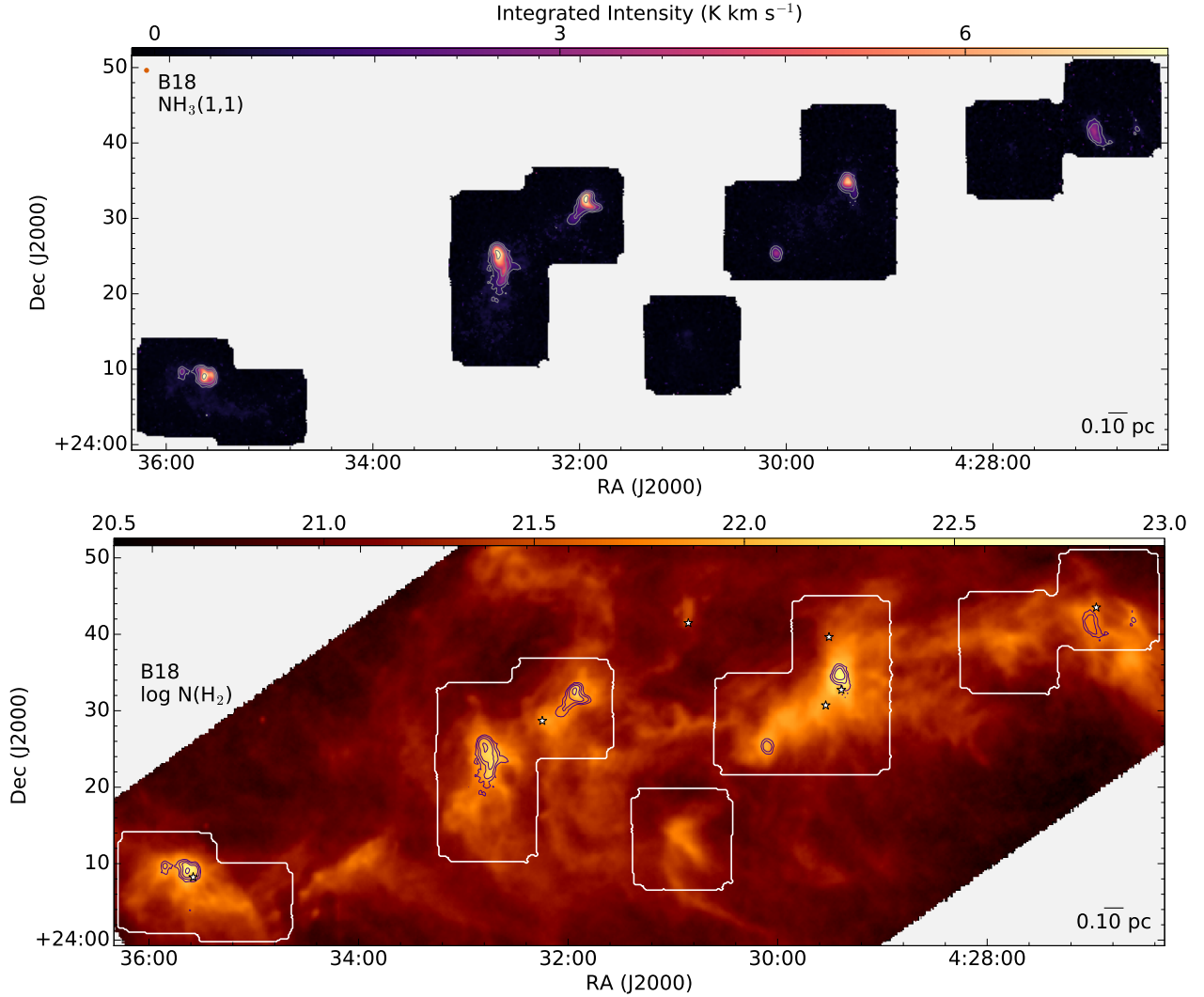


Figure 4. Top: Integrated intensity map of the NH_3 (1,1) line for the B18 region. Contours are drawn at $[3,6,12,24,\dots]\text{-}\sigma$, where σ is the rms estimated from emission-free pixels. Beam size and scale bar are shown in the top left and bottom right corners, respectively. Bottom: The H_2 column density, $\log_{10} [N(\text{H}_2)/\text{cm}^{-2}]$, derived from SED fitting of *Herschel* submillimeter dust continuum data (A. Singh et al., in preparation). Purple contours show NH_3 (1,1) integrated intensity as at top. The white contour shows the *GAS* map extent. Stars show the locations of Class 0/I and flat spectrum protostars with a reliability grade of A- or higher (Rebull et al. 2010).

system, notably of the collider species H_2 and He. The rotation temperature, T_{rot} is the temperature that characterizes the rotational level population. Our work usually focuses on the temperature that defines the population ratio between the (1,1) and (2,2) states. Finally, T_{ex} is the radiative excitation temperature of the observed spectral lines.

NH_3 fitting usually proceeds under the assumption that the (1,1) and the (2,2) line have the same radiative excitation temperature and that the frequencies are approximately equal and $\ll kT_{\text{ex}}/h$. The former assumption is only approximately true: using RADEX (van der Tak et al. 2007) for $n = 10^4 \text{ cm}^{-3}$ and $T_K = 15 \text{ K}$ and $N(\text{p} - \text{NH}_3) = 10^{14} \text{ cm}^{-2}$ gives $T_{\text{ex},(1,1)} = 8.5 \text{ K}$

and $T_{\text{ex},(2,2)} = 6.9 \text{ K}$. Each of the inversion transitions consists of multiple, resolved hyperfine components and we further assume a constant T_{ex} for all these transitions, which precludes our model representing hyperfine anomalies (e.g., Stutzki et al. 1984). We see, however, no evidence in our survey that these effects are important.

Our observations consist of antenna temperature as a function of frequency, and so we generate a model of the spectral line given the input parameters (N , T_K , T_{ex} , σ_v , v_{LSR}). We assume a single radiative excitation temperature, T_{ex} , characterizes both the (1,1) and the (2,2) transitions. We generate a model of optical depth as a function of frequency $\tau(\nu)$ to produce a model spectrum of the main-beam radiation temperature (assuming the

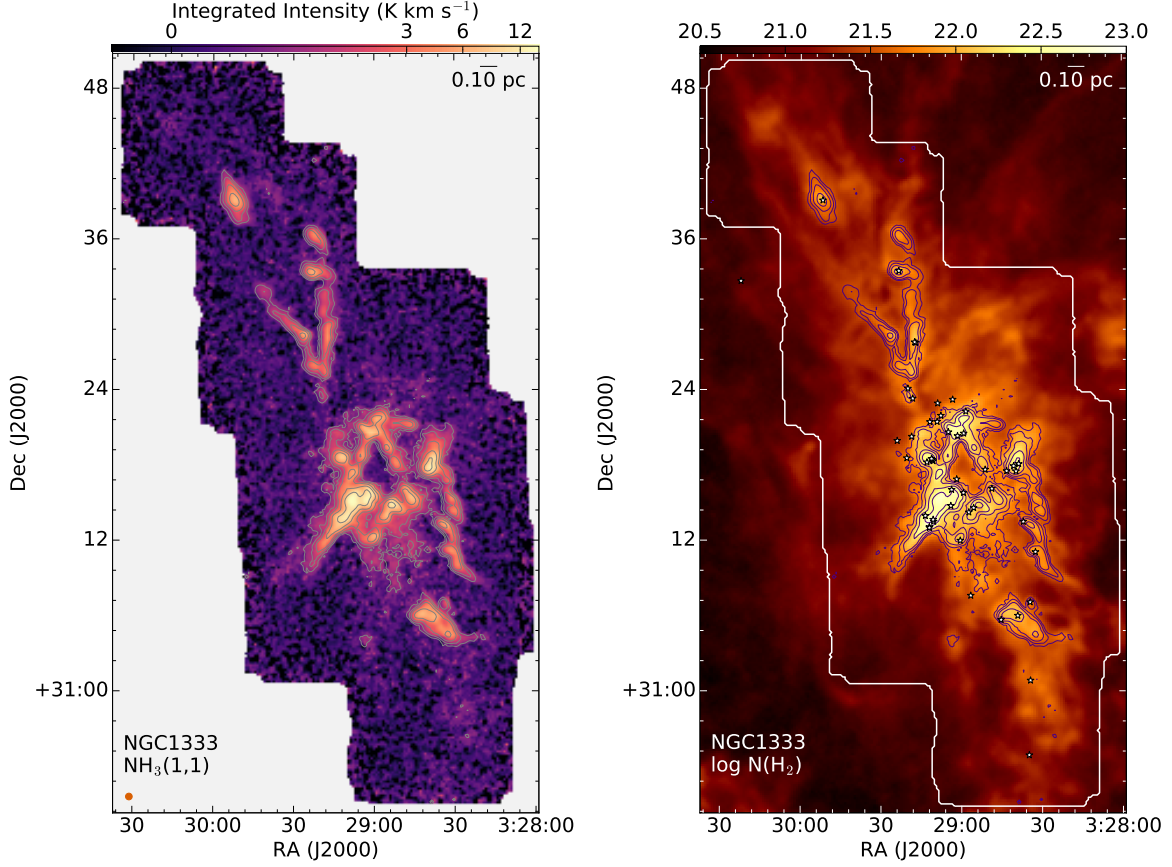


Figure 5. Like Figure 4, but for NGC 1333. Stars show the locations of Class 0/I and flat spectrum protostars (Dunham et al. 2015).

beam filling fraction = 1):

$$T_{\text{MB}}(\nu) = [J(T_{\text{ex}}) - J(T_{\text{bg}})] [1 - e^{-\tau(\nu)}]. \quad (2)$$

The optical depth, $\tau(\nu)$, is a function of the input parameters. To derive $\tau(\nu)$ from the input parameters, we first approximate the molecular population as having most of the p-NH₃ in the (1,1) and (2,2) states so that we can define a T_{rot} by the ratio of the column densities in the (2,2) and (1,1) states:

$$N_{(2,2)} = N_{(1,1)} \frac{5}{3} \exp\left(-\frac{\Delta E}{kT_{\text{rot}}}\right) \quad (3)$$

where the 5/3 is the ratio of the rotational statistical weights of the states ($g_J = 2J + 1$), and $N(J, K)$ is the sum of the column densities of the two inversion levels. From here we have that T_{rot} can be related to the kinetic temperature by analytically solving the equations of detailed balance (Swift et al. 2005):

$$T_{\text{rot}} = T_K \left\{ 1 + \frac{T_K}{T_0} \ln \left[1 + \frac{3}{5} \exp\left(-\frac{15.7 \text{ K}}{T_K}\right) \right] \right\}. \quad (4)$$

Here $T_0 = 41.5 \text{ K}$ is the energy difference between the (2,2) and the (1,1) states. We then use this rotational

temperature to calculate a partition function for the p-NH₃ molecule, though the next metastable level is the (4,4) line at 178 K above the ground state and is negligibly populated ($\tau < 1$) for $T_K < 65 \text{ K}$. The partition function is given by

$$Z_{\text{para}} = \sum (2J + 1) \times \exp\left(\frac{-h(B_{\text{rot}}J(J+1) + (C_{\text{rot}} - B_{\text{rot}})J^2)}{k_B T_{\text{rot}}}\right), \quad (5)$$

where B_{rot} and C_{rot} are the rotational constants of the ammonia molecule, 298117 and 186726 MHz, respectively (Pickett et al. 1998).

Since the excitation temperature defines the population ratio of the lower to the upper state, we can calculate N_l from the total column density in the (1,1) state, $N_{(1,1)}$.

$$\frac{N_u}{N_l} = \frac{g_u}{g_l} \exp\left(-\frac{h\nu_0}{kT_{\text{ex}}}\right), \quad (6)$$

so

$$N_l = \frac{N_{(1,1)}}{1 + \frac{g_u}{g_l} \exp\left(-\frac{h\nu_0}{kT_{\text{ex}}}\right)}. \quad (7)$$

Given the column density in the lower state, we then calculate a total, integrated optical depth in the (1,1)

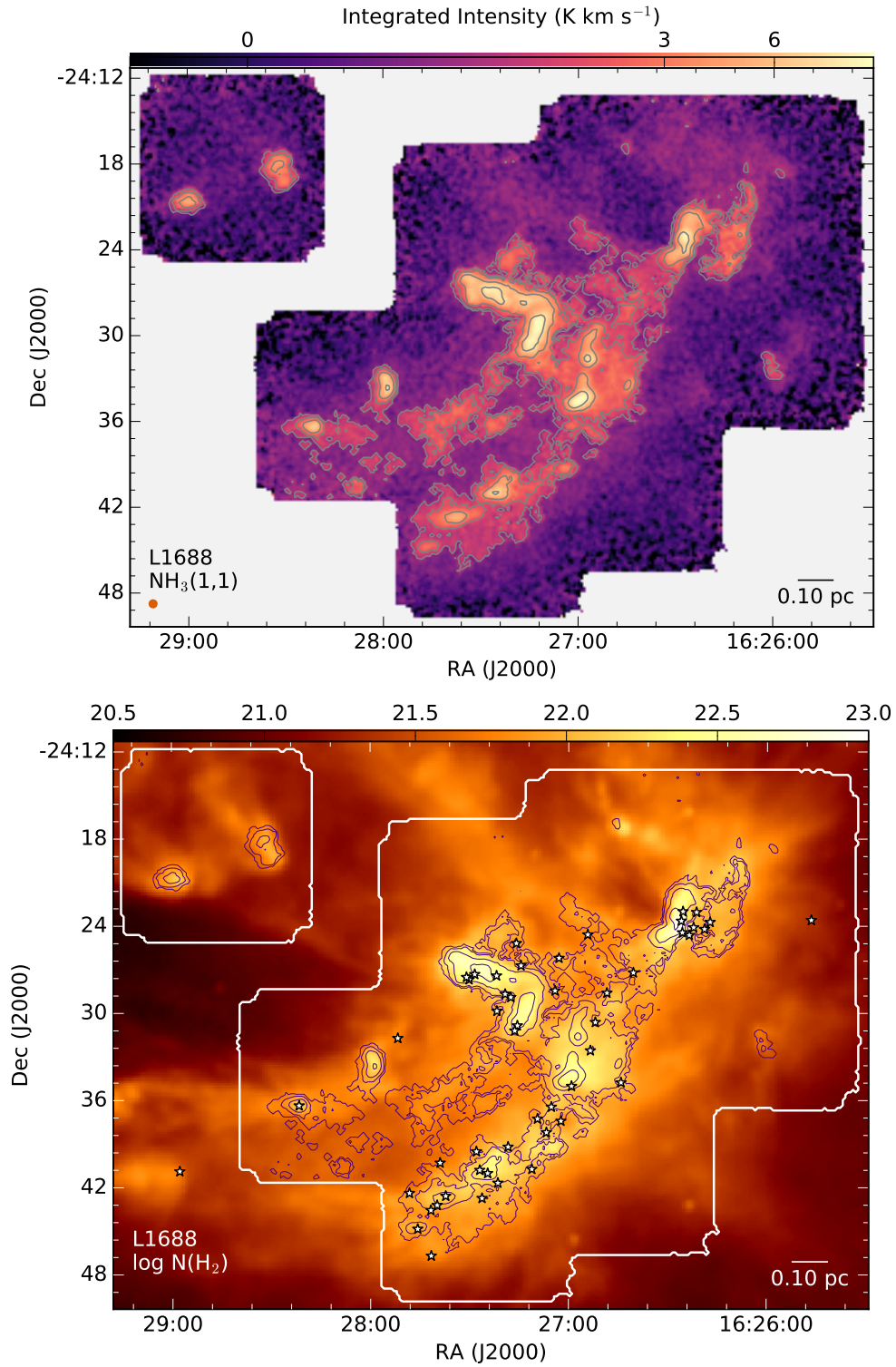


Figure 6. Like Figure 4, but for L1688. Stars show the locations of Class 0/I and flat spectrum protostars (Dunham et al. 2015).

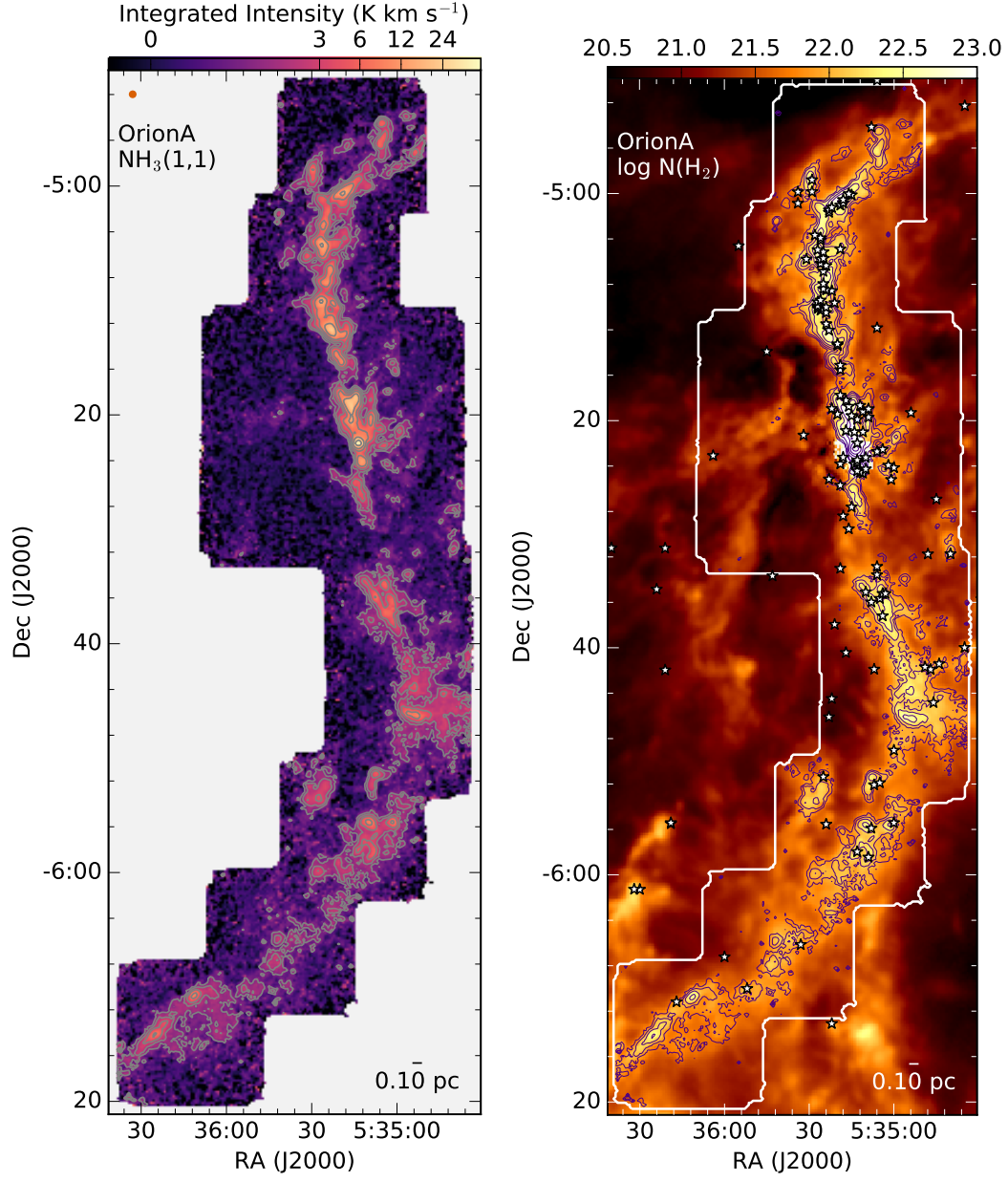


Figure 7. Like Figure 4, but for Orion A (North). Stars show the locations of Class 0/I and flat spectrum protostars (Megeath et al. 2012). Note that the column density is poorly fit toward the Orion nebula.

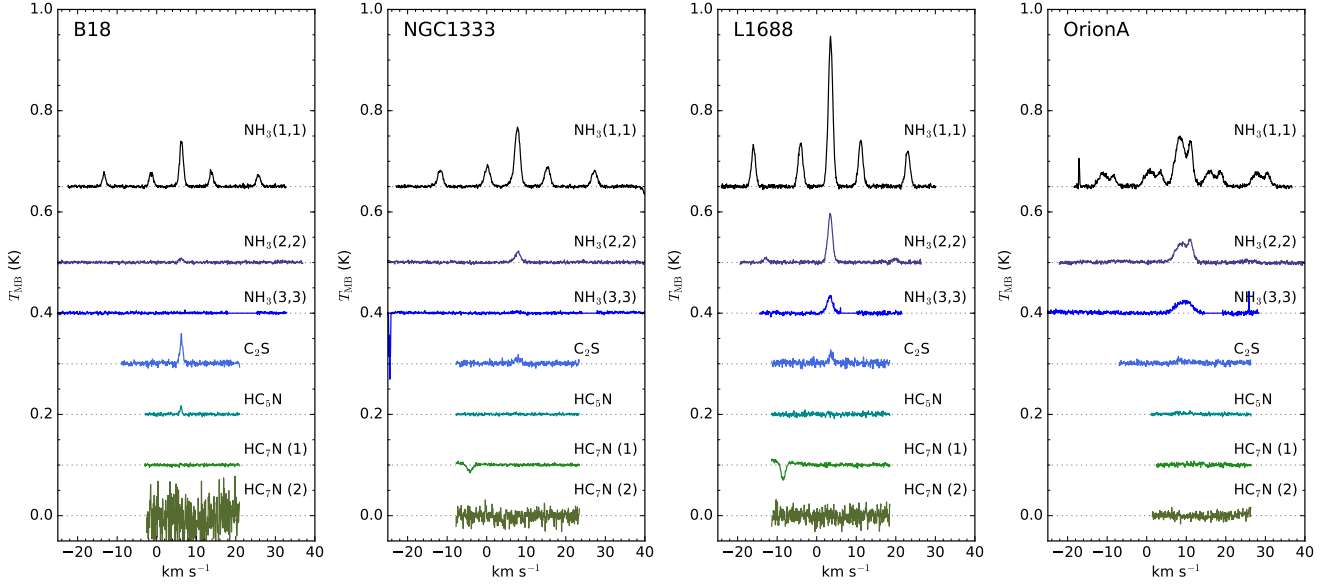


Figure 8. Spectra of all lines observed, averaged over each of the DR1 regions. Lines have been shifted in T_{MB} for clarity. The spectra labeled HC7N (1) and HC7N (2) are HC7N 21-20 and HC7N 22-21, respectively. The hyperfine structure of the NH₃ (1,1) is clearly detected in the averaged spectra toward all regions. In L1688, the relatively weaker hyperfine structure of the NH₃ (2,2) line is also seen. The HC7N 22-21 spectra have significantly more noise than the other lines, primarily due to greater rms values at the maps' edges, and have been multiplied by a factor of 0.5. In NGC 1333 and L1688, the absorption features in the HC7N 21-20 spectra are artifacts of frequency-switching while observing, where some of the NH₃ (1,1) emission was switched near the HC7N line. After running the calibration and imaging pipelines, several footprints have greater rms noise values at some velocity channels. For NH₃ (3,3), we have masked the velocity channels where this issue is strongest, but several additional features (visible as sharp, narrow spikes) remain.

and (2,2) lines using standard radiation relationships. In particular, the opacity per unit frequency,

$$\alpha_\nu = \frac{c^2}{8\pi} \frac{1}{\nu_0^2} \frac{g_u}{g_l} n_l A_{ul} \left[1 - \exp\left(-\frac{h\nu_0}{kT_{\text{ex}}}\right) \right] \phi(\nu), \quad (8)$$

where ν_0 is the line rest frequency, n_l is the volume density in the lower state of the transition, g_u, g_l are statistical weights of the upper and lower states of a transition and whose ratio is unity for each pair of inversion transitions, A_{ul} is the Einstein A coefficient, and $\phi(\nu)$ is the line profile function defined such that

$$\int \phi(\nu) d\nu = 1. \quad (9)$$

We integrate over line of sight and frequency to get the optical depth as a function of the column density in the lower transition of the spectral line:

$$\begin{aligned} \tau &= \int \alpha_\nu ds d\nu \\ &= \frac{c^2}{8\pi} \frac{1}{\nu_0^2} \frac{g_u}{g_l} N_l A_{ul} \left[1 - \exp\left(-\frac{h\nu_0}{kT_{\text{ex}}}\right) \right]. \end{aligned} \quad (10)$$

or, in terms of $N_{(1,1)}$, which is the total column of molecules in the symmetric and antisymmetric (upper

and lower) states of the J=1 K=1 level,

$$\begin{aligned} \tau_{(1,1)} &= \int \alpha_\nu ds d\nu \\ &= \frac{c^2}{8\pi} \frac{1}{\nu_0^2} \frac{g_u}{g_l} N_{(1,1)} A_{ul} \frac{1 - \exp\left(-\frac{h\nu_0}{kT_{\text{ex}}}\right)}{1 + \frac{g_u}{g_l} \exp\left(-\frac{h\nu_0}{kT_{\text{ex}}}\right)}. \end{aligned} \quad (11)$$

The optical depth for the (2,2) transition has the same form. The column densities are related to the total column N through the partition function via T_{rot} .

As per Mangum & Shirley (2015),

$$A_{ul} = \frac{64\pi^4 \nu_0^3}{3hc^3} \frac{1}{g_u} |\mu_{ul}|^2, \quad (12)$$

where μ_{ul} is the dipole moment of a given transition, which is related to the dipole moment of the molecule μ^2 by

$$|\mu_{ul}|^2 = \mu^2 \frac{K^2}{J(J+1)}, \quad (13)$$

where $\mu = 1.468 \times 10^{-18}$ esu cm for ammonia.

Finally, we calculate the line profile function $\phi(\nu)$ from the hyperfine structure of the different transitions. Since each hyperfine transition is assumed to have the same

radiative excitation temperature, $\phi(\nu)$ is given as the weighted sum of Gaussian line profiles:

$$\phi(\nu) = \sum_i \frac{w_i}{\sqrt{2\pi}\sigma_{\nu,i}} \exp \left[-\frac{(\nu - \nu_0 - \delta\nu_i)^2}{2\sigma_{\nu,i}^2} \right]. \quad (14)$$

The weights, w_i , and frequency offsets from the rest frequency, $\delta\nu_i$, are set by quantum mechanics and are tabulated in [Mangum & Shirley \(2015\)](#). The frequency widths of the lines are calculated in terms of the velocity widths from the Doppler formula: $\sigma_{\nu,i} = (\nu_0 + \delta\nu_i)\sigma_v c^{-1}$.

This model provides a complete spectrum T_{MB} as a function of input parameters. We then use a non-linear gradient descent algorithm (MPFIT; [Markwardt 2009](#)) implemented in Python to calculate the optimal values of the input parameters and their uncertainties (e.g., [Press et al. 2002](#)).

$$\begin{aligned} & \text{argmin } \chi^2 = \\ & \text{argmin } \left[\sum_j \frac{[T_{\text{MB,obs}}(\nu_j) - T_{\text{MB}}(\nu_j; N, T_K, T_{\text{ex}}, \sigma_v, v_{\text{LSR}})]^2}{\sigma^2(\nu_j)} \right] \end{aligned} \quad (15)$$

Non-linear least squares fitting can be unstable unless the algorithm is provided with a good set of initial conditions. In particular, the v_{LSR} value of the fit is critical, so we use guesses based on the velocity centroid of the line (or of the center hyperfine component for ammonia lines). The velocity dispersion is taken initially to be the intensity weighted second moment of the velocity around the centroid. We adopt $\log_{10}(N/\text{cm}^{-2}) = 14.5$, $T_K = 12$ K, and $T_{\text{ex}} = 3$ K as our initial guesses. The fit is performed serially for each position in the map, starting from the pixel with the highest integrated intensity and working down in reverse order of integrated intensity. If a given position has a neighboring pixel with a good fit, the parameters of the successful fit are used as initial guesses for the next optimization.

In this release, we do not attempt to fit multiple velocity components or non-Gaussian velocity profiles to our spectra. Such approaches require human supervision under current implementations ([Henshaw et al. 2016b](#)). Our fits produce good quality results for most of the studied regions ($> 95\%$ of the detections in all regions release in DR1), but some lines of sight in NGC 1333, L1688, and Orion A show good evidence for needing improved models for their underlying velocity components.

Recent work by [Estalella \(2017\)](#) presented the HFS hyperfine fitting tool. The assumptions embedded in our approach are essentially equivalent to those embedded in that tool, though [Estalella \(2017\)](#) used a relationship between T_{rot} and T_K based on the collision coefficients in [Maret et al. \(2009\)](#), which differs from the [Swift et al. \(2005\)](#) expression by $> 10\%$ at $T > 50$ K. The optimization approach differs significantly, in that we use a gradient descent non-linear regression, whereas

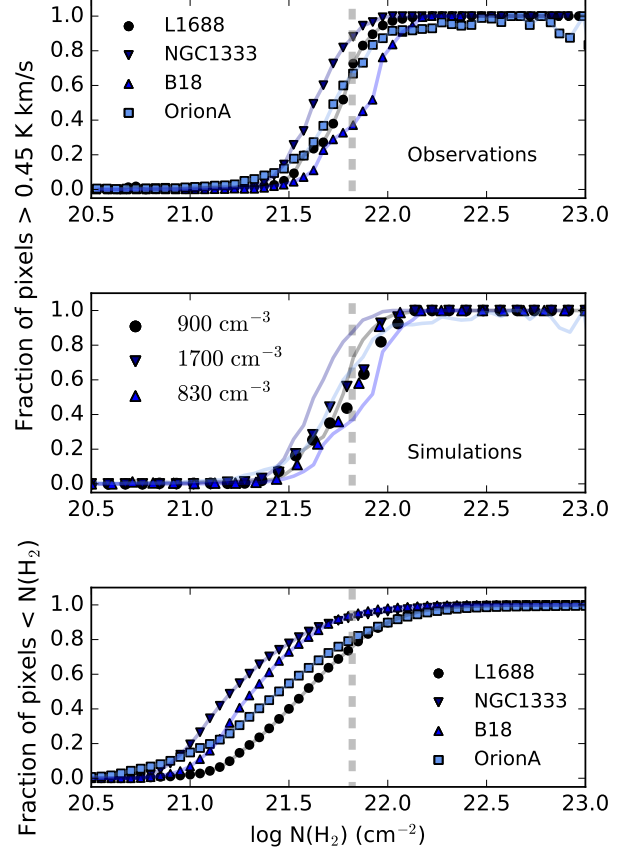


Figure 9. Top: The fraction of pixels with NH_3 (1,1) integrated intensity $\geq 0.45 \text{ K km s}^{-1}$ (approximately 3σ) in each DR1 region. We show only the range in $N(\text{H}_2)$ where significant change in NH_3 detections are seen; the largest $\log N(\text{H}_2)$ values reach ~ 26 in Orion A. In all panels, the vertical dashed grey line shows $A_V \simeq 7$. Middle: Fraction of NH_3 pixels as a function of $N(\text{H}_2)$ for the GAS data (lines; colors as at top) and a synthetic observation (markers) based on astrochemical modeling from [Offner et al. \(2014\)](#). The simulated values are obtained from maps with average number densities of $n_{\text{H}_2} = 900 \text{ cm}^{-3}$, 1700 cm^{-3} , and 830 cm^{-3} . The regions have projected sizes of 2 pc, 0.5 pc, and 1 pc, respectively. The synthetic observation assumes a circular beam of $32''$ and a distance of 250 pc with a $\Delta v = 0.08 \text{ km s}^{-1}$ velocity channel width. Bottom: The cumulative distribution function of $N(\text{H}_2)$ within the observed footprints.

the HFS code uses a genetic algorithm. Our approach requires more accurate guesses for the initial conditions, but the strategy we have adopted leads to good convergence. The uncertainty estimates both follow standard estimates based on χ^2 fitting.

3.2. v_{LSR} -tracked integrated intensity maps

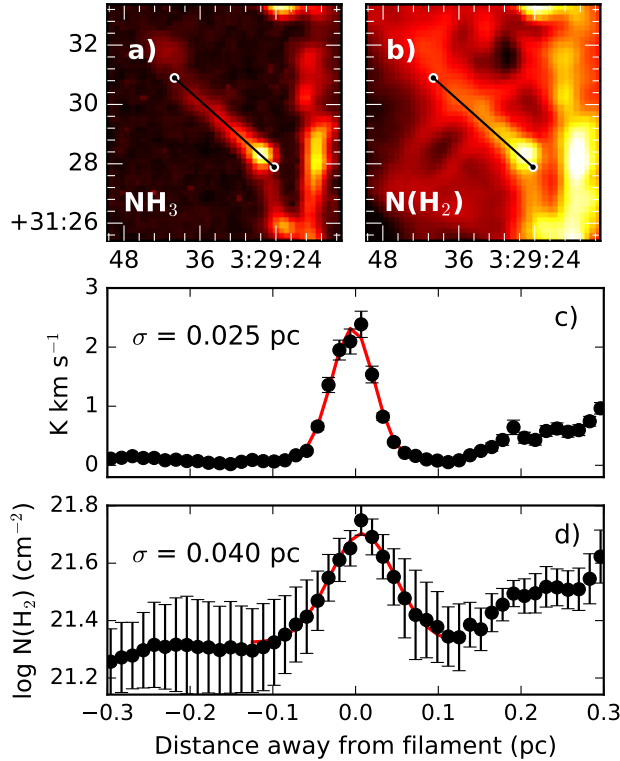


Figure 10. a) NH_3 (1,1) integrated intensity toward a filament in NGC 1333. b) $N(\text{H}_2)$ over the same region as in a). c) The radial profile of the filament in NH_3 (1,1) integrated intensity, averaged along the straight black line identified in a) and b). d) The radial profile of the filament in $N(\text{H}_2)$ along the same line.

When calculating the integrated intensity map of a specific line, it is common practice to define a fixed velocity range that includes the emission of all hyperfine components for a single region. The maps presented here, however, cover large areas of individual star-forming clouds. In these regions, large velocity gradients make a fixed range sub-optimal.

To create the NH_3 (1,1) and (2,2) integrated intensity maps, we therefore first produce a model emission cube given the best fit models over the entire map. We set a threshold value > 0.0125 K (i.e., slightly greater than machine precision) to identify the spectral channels with model line emission. We then select all the voxels (3D “pixels”) with brightness greater than the threshold to calculate the integrated intensity map. In regions where no line fit was found, we calculate the integrated intensity within a spectral window defined by the mean cloud v_{LSR} and σ_v . We calculate the uncertainty at each pixel based on the channel range thus identified.

This method for calculating the integrated intensity while following the velocity gradients in a star-forming region depends on having a good model, which is usually

the case for NH_3 emission in the clouds presented here. There are locations, however, where the emission is not perfectly fitted with a single component. In those cases, this method may not include some of the flux in the final moment maps. A detailed study of the gas kinematics will be done in future papers focusing on individual regions.

Since the spatial and kinematic distribution of NH_3 (3,3) and the carbon-chain molecules do not typically follow that of NH_3 (1,1) and (2,2), and are significantly less extended, where detected we use a single velocity range to produce integrated intensity maps of these transitions.

3.3. Data masking

For all regions, NH_3 (1,1) and (2,2) lines were fit where the NH_3 (1,1) line had a $\text{SNR} > 3$ through the property map analysis described above. Here $\text{SNR} = T_{\text{peak}}/\text{rms}$, where T_{peak} is the peak brightness temperature in the NH_3 (1,1) cube at a given pixel. At this SNR, however, not all line parameters could be fit well over all pixels. We consequently mask the output parameter maps to remove poor fits to the NH_3 lines. Since v_{LSR} , σ_v , and T_{ex} require good fits to the NH_3 (1,1) line only, we mask pixels with a $\text{SNR} < 3$ in NH_3 (1,1) integrated intensity. The parameters T_{K} and N require detections of the NH_3 (2,2) line, and for these property maps we additionally mask any pixels with a $\text{SNR} < 3$ in the peak NH_3 (2,2) line intensity. Finally, we apply a mask that removes any remaining isolated pixels that are unconnected with larger-scale features and are generally noisy datapoints. Additional masking based on the returned uncertainty in the fit parameters may be done for detailed analysis in future papers.

4. RESULTS

4.1. NH_3 moment maps

In Figures 4, 5, 6, and 7, we show the NH_3 (1,1) integrated intensity maps of B18, L1688, NGC 1333, and Orion A (North). We further overlay the NH_3 (1,1) integrated intensity contours on maps of the H_2 column density, $N(\text{H}_2)$, in each region, derived from spectral energy distribution (SED) modeling of continuum emission from dust observed with the *Herschel Space Observatory* (A. Singh et al., in preparation). The SED modeling is described in more detail in § 4.3. The locations of Class 0/I and flat spectrum protostars are highlighted, based on infrared analysis of Taurus (Rebull et al. 2010), Perseus and Ophiuchus (Dunham et al. 2015), and Orion (Megeath et al. 2012).

In general, the NH_3 (1,1) integrated intensity emission follows closely the H_2 column density as traced by dust continuum emission. The GAS data highlight individual cores and filamentary structure, as expected, but at the achieved sensitivity level also reveal low-SNR emission that extends between the higher column density structures, and follows, by visual inspection, the

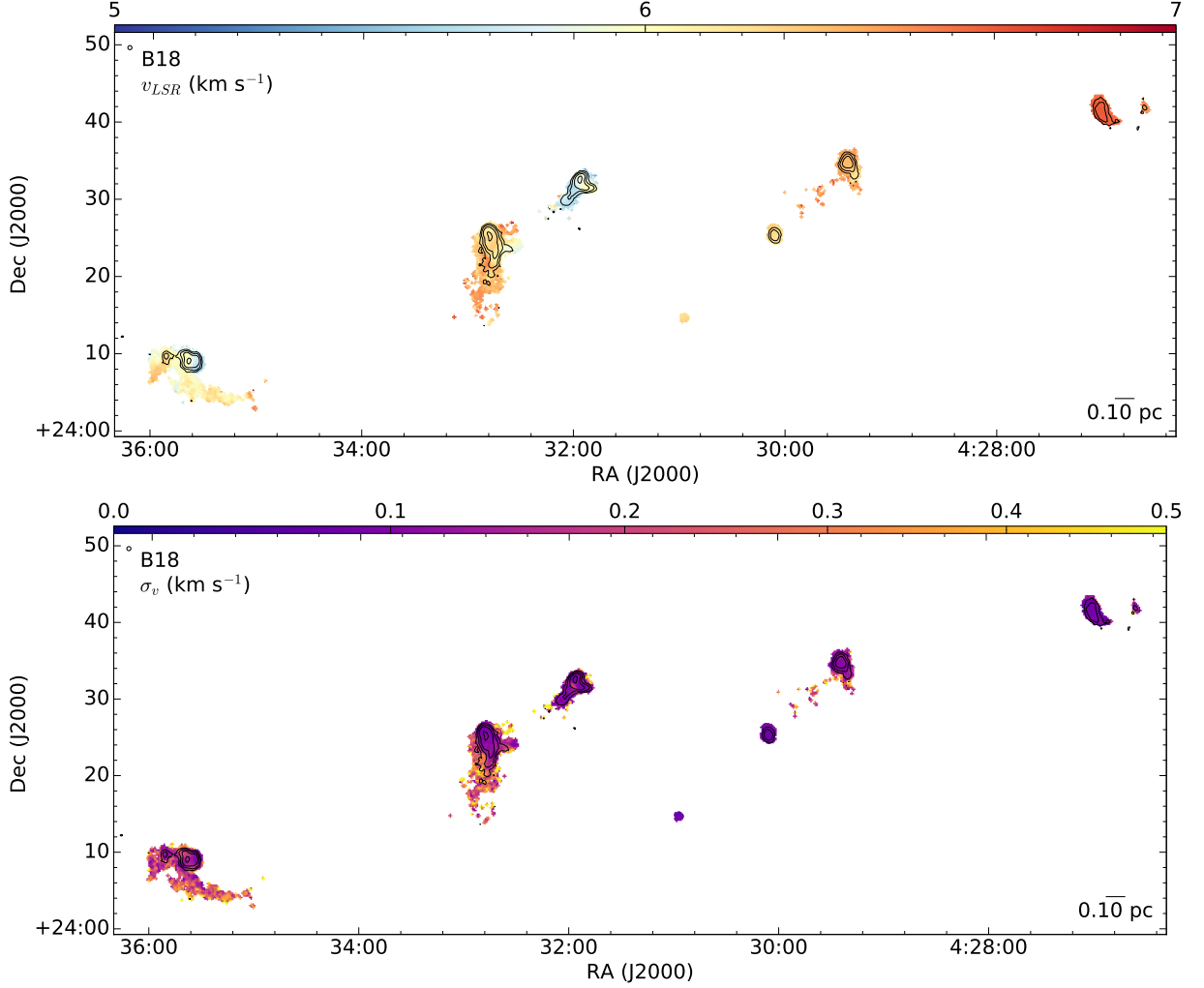


Figure 11. Top: v_{LSR} (km s^{-1}) in B18. Contours are NH_3 (1,1) integrated intensity as in Figure 4. Bottom: σ_v (km s^{-1}) in B18, with the color scale chosen to match the spread in σ_v across all DR1 regions.

lower H_2 column density material. As an example, Figure 4 shows how NH_3 (1,1) integrated intensity peaks highlight the dense cores in Taurus B18, while extended emission matches the shape of the surrounding material seen in $N(\text{H}_2)$. Toward L1688, the central, high column density cores are surrounded by material at lower $N(\text{H}_2)$ values that fade away toward the northeast in Figure 6. Low SNR NH_3 (1,1) emission is similarly more extended in this direction. Extended, low SNR NH_3 emission can also be seen toward the more distant NGC 1333 and Orion A regions.

We show in Appendix B the integrated intensity maps of NH_3 (2,2) and NH_3 (3,3) (where detected) toward the DR1 regions.

In appearance, the NH_3 (2,2) integrated intensity maps match closely the NH_3 (1,1) maps, but with lower SNR.

In NGC 1333, L1688, and Orion A, NH_3 (3,3) emission is detected, indicative of warmer dense gas present

in these regions. In NGC 1333, the NH_3 (3,3) emission is compact and centered on two well-known sources, the bright Class 0/I object SVS 13 and Herbig-Haro object HH12. In L1688, the NH_3 (3,3) emission is more diffuse, overlapping the NH_3 (1,1) contours in the northwest but primarily visible on the western edge. This region is an interface between the cold, dense gas of the Ophiuchus A core and an edge-on photodissociation region associated with the B2V star HD 147889 (Loren & Wootten 1986; Liseau et al. 1999). The NH_3 (3,3) emission in Orion A peaks at the BN/KL region and traces the associated filamentary structure extending north. The NH_3 (3,3) emission is also prominent toward the Orion bar, highlighting this feature significantly better than NH_3 (1,1).

4.2. Detections of other lines

In the previous sections, we have focused primarily on emission from the NH_3 transitions toward the DR1 regions. We additionally mapped the rotational transi-

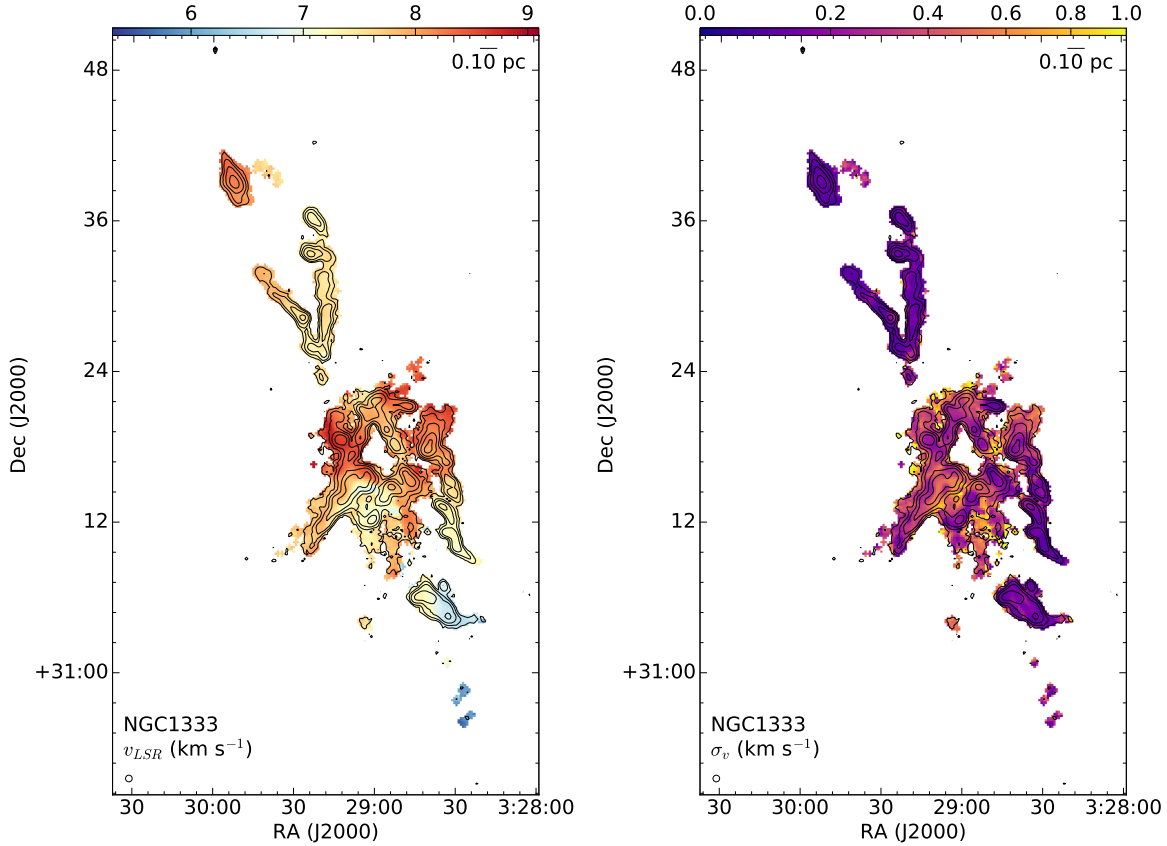


Figure 12. Like Figure 11 but for NGC 1333.

tions of several carbon chains, which are listed in Table 2. Not all lines were detected in all regions, and the HC_7N transitions were not detected above the rms noise or in an integrated intensity map in any region. We show in Appendix C integrated intensity maps of the additional detected lines in each region. In Figure 8, we show the spectrum of each observed line averaged over each DR1 region after masking of the edge pixels as described in § 2.5.

Toward B18, both HC_5N 9-8 and C_2S $2_1 - 1_0$ are detected. For both molecules, the distribution of emission is significantly offset from NH_3 (1,1), generally peaking away from the NH_3 emission peaks. While HC_5N is seen interior to some NH_3 (1,1) contours, C_2S , where seen, appears to surround the NH_3 emission. Similar distributions of NH_3 and C_2S have been identified toward the L1495-B218 filaments in Taurus (Seo et al., in preparation). This offset between emission from NH_3 and carbon-chain molecules is common in nearby star-forming regions and is generally explained by the depletion of C-bearing molecules onto dust grains at high densities, while NH_3 remains in the gas phase. Asymmetric offsets like those seen in B18 might be indicative of the contraction of dense gas (Tafalla et al. 2004) or accretion of lower density material onto the dense cores (Friesen et al. 2013).

HC_5N is also detected toward NGC 1333 in a single compact region (slightly greater in extent than the beam) located to the west of the central structure traced by NH_3 (1,1). While C_2S is not visible in the integrated intensity map of NGC 1333 (not shown), the averaged C_2S spectrum in Figure 8 shows a significant detection. Similarly, C_2S is clearly detected toward L1688 in the averaged spectrum, while the integrated intensity map shows only a faint, $\text{SNR} \sim 3 - 4$ feature near the map center. In future analysis we will investigate extended, low-level emission of NH_3 and the other observed lines through spectral stacking techniques.

4.3. Relationship of NH_3 emission to $N(\text{H}_2)$

4.3.1. Column density of H_2

The H_2 column density, $N(\text{H}_2)$, used here and in the maps above was determined in two steps (A. Singh et al., in preparation). First, the optical depth of dust was obtained by fitting a modified blackbody to the spectral energy distribution (SED) of the thermal continuum emission from cold dust using *Herschel* maps at 160 μm , 250 μm , 350 μm , and 500 μm . Second, this is converted to gas column density by a simple scale factor. An additional scale factor can be used to convert to A_V .

To implement the first step, we note that for dust temperature T_d the intensity of thermal dust emission

is

$$I_\nu = (1 - e^{-\tau_\nu})B_\nu(T_d) \quad (16)$$

and we adopt $\tau_\nu = \tau_{\nu_0}(\nu/\nu_0)^\beta$. The dust opacity spectral index was taken to be $\beta = 1.62$, the mean value found by Planck over the entire sky (Planck Collaboration XI 2014). All maps were first convolved to the $36''$ resolution of the $500 \mu\text{m}$ data and put on the same grid. Zero offsets were determined for each wavelength by comparison with Planck emission maps. Minor color corrections were applied during the SED fitting. The SED parameters τ_{ν_0} and T_d were determined at each pixel using χ -squared minimization.

The required scale factor to convert to $N(\text{H}_2)$ can be seen via the relation

$$\tau_{\nu_0} = \kappa_{\nu_0} \mu m_{\text{H}} N(\text{H}_2) \quad (17)$$

where $\mu = 2.8$, m_{H} is the mass of a Hydrogen atom, and κ_{ν_0} is the opacity (incorporating a gas-to-dust mass ratio of 100), which is taken to be $0.1 \text{ cm}^2 \text{ g}^{-1}$ at $\nu_0 = 1000 \text{ GHz}$ (Hildebrand 1983). The considerable systematic uncertainty in the opacity propagates uniformly throughout the map into a systematic uncertainty in the scale of $N(\text{H}_2)$.

Ideally, one could scale the dust-based τ_{ν_0} directly into A_V , which is also proportional to dust column density. This scaling too is uncertain (Planck Collaboration XI 2014). Here we simply use $N(\text{H}_2) = 9.4 \times 10^{20} \text{ cm}^{-2}$ (A_V/mag), which is based on data for the diffuse interstellar medium, not molecular clouds, and so the two-figure precision given belies the actual uncertainty.

Regional variations in dust spectral properties could impact the values determined through SED modeling (and thus the threshold for NH_3 emission discussed below), although this is likely a small effect compared to the uncertainty in the scaling: for example, a change in β of 0.05 produces a change in $N(\text{H}_2)$ of only $\sim 5 \%$.

4.3.2. $N(\text{H}_2)$ threshold for NH_3 emission

Figure 9 (top) shows a smooth increase in the fraction of pixels with NH_3 emission as $N(\text{H}_2)$ increases, with some variation between regions in the rate of rise in the correlation, and in the H_2 column density where NH_3 is found universally. The survey initially targeted regions with extinctions $A_V > 7$, based on previous observations of NH_3 in nearby star-forming regions. That this was a good choice is confirmed by the DR1 data, where above this extinction value (highlighted by the gray dashed line in the Figure, assuming $N(\text{H}_2) = 9.4 \times 10^{20} \text{ cm}^{-2}$ (A_V/mag)), $\gtrsim 60 \%$ of pixels have detectable NH_3 emission above the survey noise levels in all DR1 regions, with the exception of B18, where the threshold appears to be slightly higher. Furthermore, the survey footprints were rectangular, not following the high column density contours precisely, and so lines of sight corresponding to lower column density were surveyed too.

4.3.3. Comparison to simulations

We compared the GAS data to results calculated from a synthetic observation of NH_3 emission produced by a 3D hydrodynamic simulation. The simulated emission map is based on astrochemical modeling performed by Offner et al. (2014). The abundances were computed by post-processing the simulation with 3D-PDR, a three-dimensional photodissociation region code (Bisbas et al. 2012), assuming the simulation is irradiated with a uniform 1 Draine field. We next computed the emission using RADMC-3D⁴. Finally, to model instrumental affects we convolved the emission map with a $32''$ beam for a distance of 250 pc and added Gaussian noise with $\sigma_{\text{rms}} = 0.15 \text{ K km s}^{-1}$, a typical value from the GAS NH_3 (1,1) integrated intensity maps.

Figure 9 (middle) shows the emission fraction for the total simulation domain ($n_{\text{H}_2} = 900 \text{ cm}^{-3}$) and for two sub-regions ($n_{\text{H}_2} = 1,700 \text{ cm}^{-3}$ and $n_{\text{H}_2} = 830 \text{ cm}^{-3}$). These distributions agree well with the GAS distributions and demonstrate that variation in the underlying gas densities can account for some of the spread between regions. Differences between regions may also result from variations in the local UV radiation field and mean column density, which we will explore in future work.

Lastly, Figure 9 (bottom) shows the cumulative distribution of $N(\text{H}_2)$ in the observed footprints of the GAS DR1 regions. A more gradual increase is seen relative to the top panel, where the presence of NH_3 rises steeply above $\log_{10} N(\text{H}_2)/[\text{cm}^{-2}] \sim 21.5$.

Figure 9 shows that observations of NH_3 inversion transitions become an excellent indicator of where there is gas at higher column densities, and thus probably higher volume densities, but miss material at lower column density that can be detected via the thermal dust continuum emission.

4.3.4. The appearance of filaments

The ability of NH_3 to highlight higher density structures enables analyses of structure sizes, masses and concentrations that are complementary to similar analyses using dust continuum emission. For example, we noted previously that filaments are prevalent within star-forming regions, with a typical filament width of $\sim 0.1 \text{ pc}$ suggested by profile-fitting of the mid- and far-infrared dust emission (Arzoumanian et al. 2011). The authors define the filament ‘width’ as the FWHM of a Gaussian fit to the innermost section of the filament’s radial profile. In Figure 10, we show a filament in NGC 1333 in a) NH_3 (1,1) integrated intensity, and b) $N(\text{H}_2)$. In panels c) and d), we show the filament radial profile in NH_3 and $N(\text{H}_2)$, respectively, averaged along the long axis shown by the straight line in panels

⁴ <http://www.ita.uni-heidelberg.de/~dullemond/software/radmc-3d>

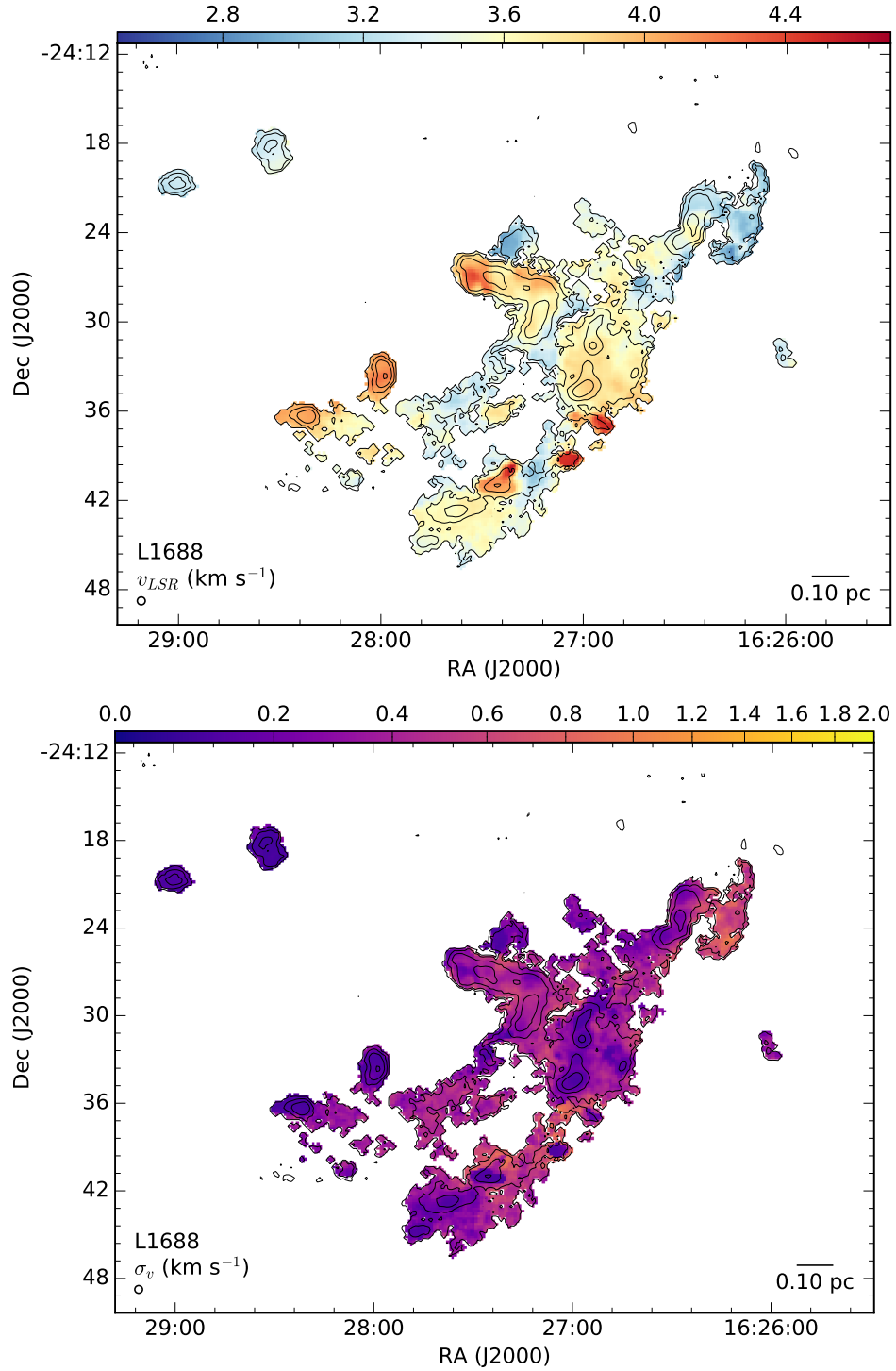


Figure 13. Like Figure 11 but for L1688.

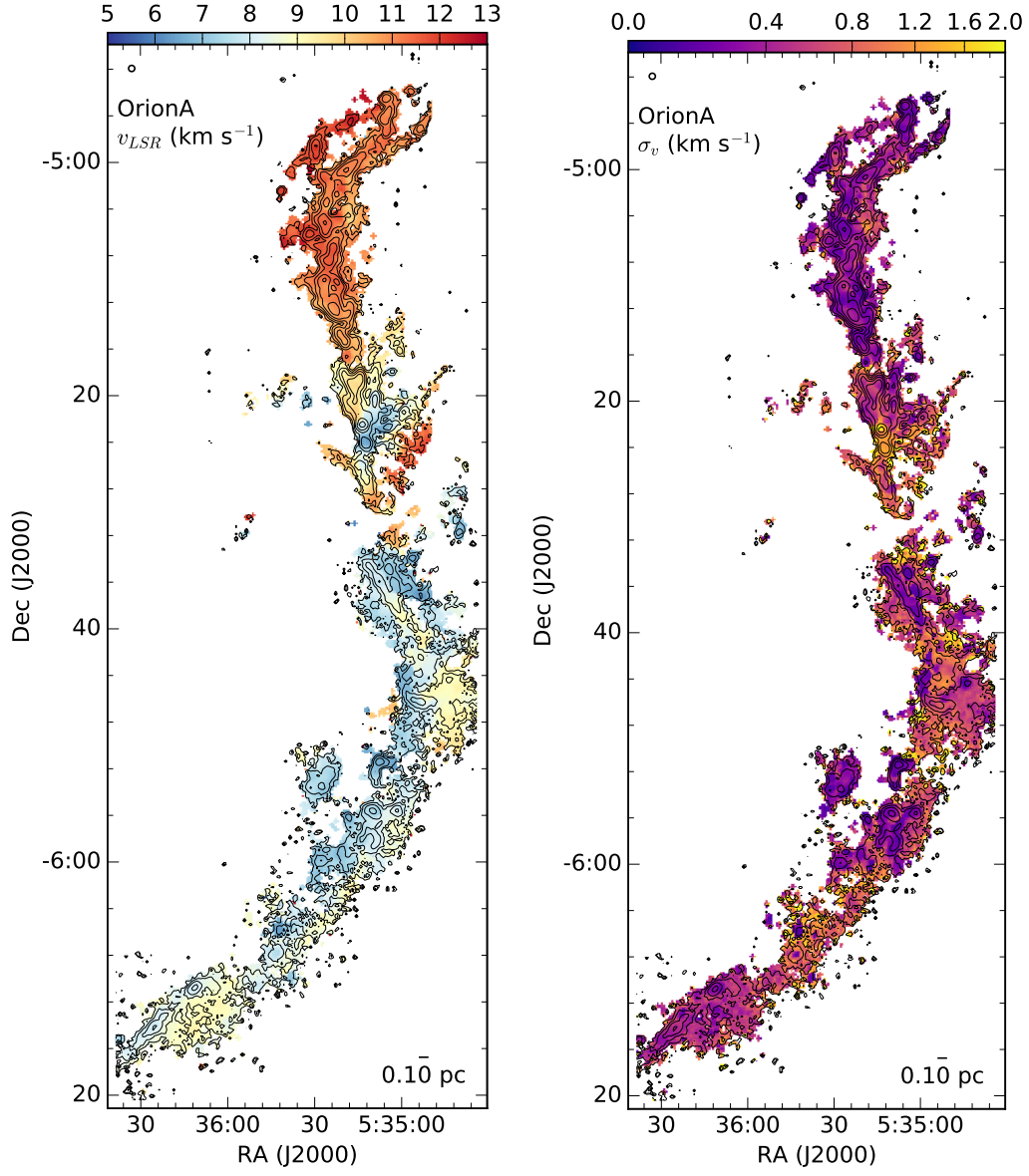


Figure 14. Like Figure 11 but for Orion A (North).

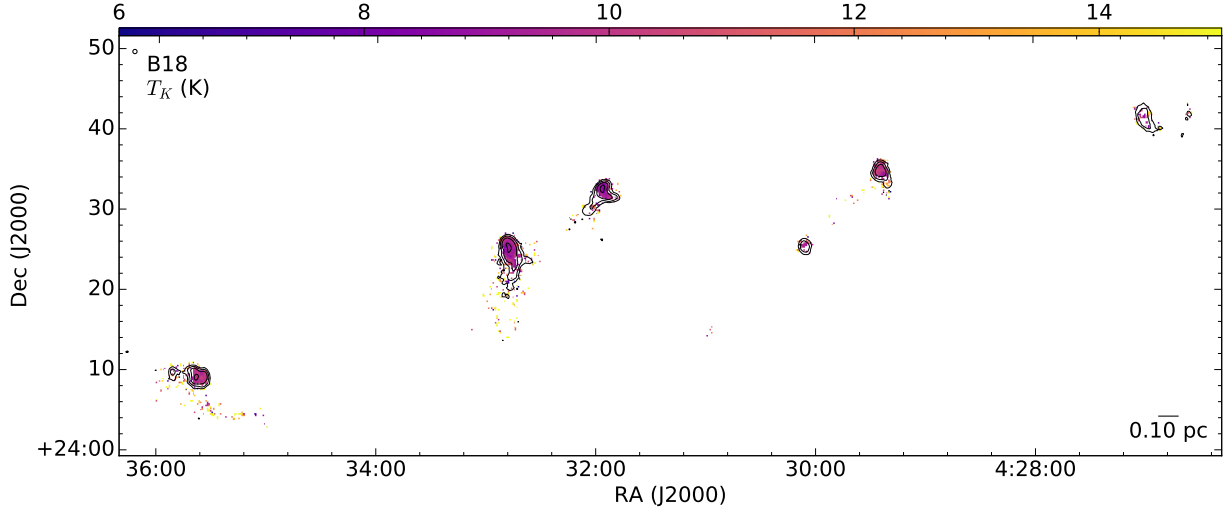


Figure 15. T_K (K) in B18, with the color scale chosen to match the spread in T_K across all DR1 regions. Contours are NH_3 (1,1) integrated intensity as in Figure 4.

a) and b). We then fit a Gaussian, plus constant offset, to the innermost 0.1 pc of each radial profile, matching the analysis of [Arzoumanian et al.](#), with the results shown in red. Similar to the authors' result, we find a filament FWHM of 0.086 pc in $N(\text{H}_2)$, but show that the filament is narrower in NH_3 emission by a factor ~ 1.4 .

Locating the range of $N(\text{H}_2)$ values from panel d) in the top panel of Figure 9 shows that the NH_3 emission does not measure the full column density, only the denser core. Furthermore, in $N(\text{H}_2)$ the filament is embedded within lower column density material and crossed by additional, narrow features, while the NH_3 emission highlights the filament only.

Future work using the GAS dataset will identify filamentary structures systematically using algorithms such as *filfinder* ([Koch & Rosolowsky 2015](#)) and *DisPerSE* ([Sousbie 2011](#)). This will probe the properties of filamentary structures at different spatial scales and determine what dominates the kinematics within the filamentary structures at each scale, thus providing insight into their origin.

4.4. Property Maps

We present the property maps resulting from the NH_3 line fitting described in § 3. In Figures 11, 12, 13, and 14, we show maps of v_{LSR} and σ_v for each of the DR1 regions. The gas temperature T_K is presented for all regions in Figures 15, 16, 17, and 18. Maps of the NH_3 column density, $N(\text{NH}_3)$, are shown for all regions in Appendix D.

While detailed analyses of the dense gas properties traced by NH_3 in the Gould Belt will be presented in future papers, we discuss here some general results and trends in the data.

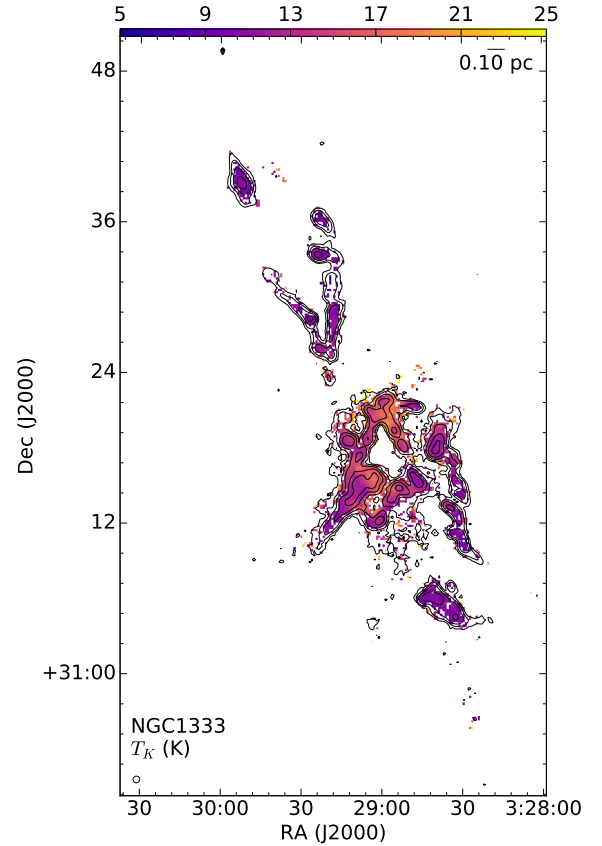


Figure 16. Like Figure 15 but for NGC 1333.

Figures 11 - 14 show significant variations in the kinematic properties traced by NH_3 between the DR1 regions. In both B18 and L1688, gas velocities are concentrated around the mean cloud value with only a small spread. In B18, at the survey sensitivity, we detect NH_3

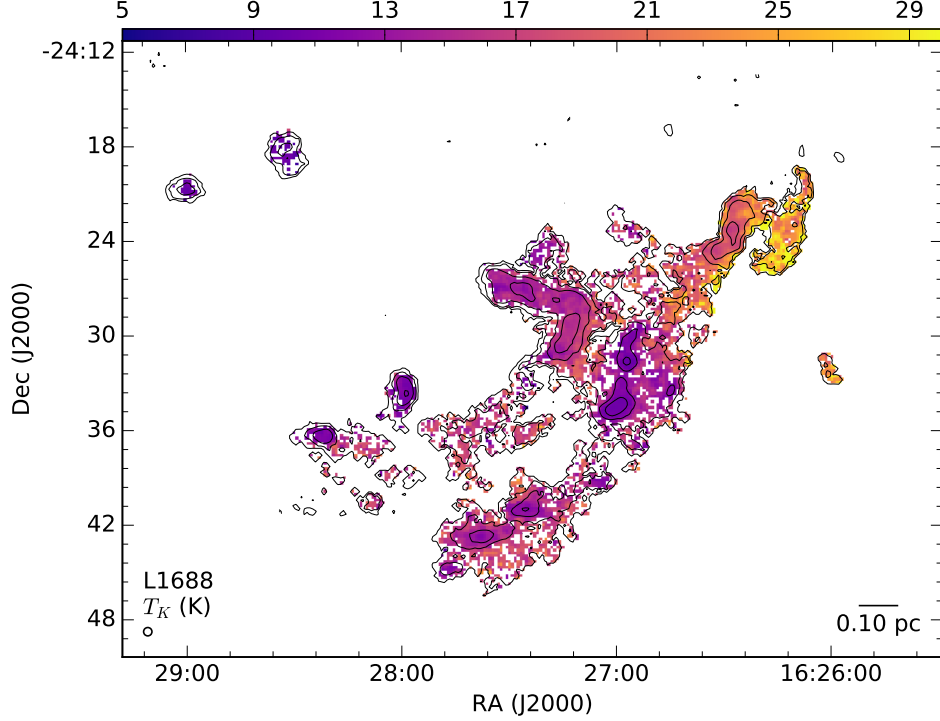


Figure 17. Like Figure 15 but for L1688.

toward only a small fraction of the extended gas surrounding the NH_3 peaks, and the v_{LSR} values shown represent mainly the v_{LSR} of the densest NH_3 structures in the region. In the other DR1 regions, we detect more of the extended NH_3 emission. The v_{LSR} map of L1688 in Figure 13 shows that distinct variations in v_{LSR} are seen between some of the NH_3 structures and the surrounding dense gas, but the velocity changes are small ($\lesssim 1 \text{ km s}^{-1}$). Toward NGC 1333, a broader distribution in v_{LSR} is visible in Figure 12. Unlike the variation in v_{LSR} between compact and extended emission in L1688, gradients in v_{LSR} are seen along the filamentary structures in NGC 1333. A large gradient in v_{LSR} is also clearly visible extending along the Orion A filament from north to south, with a smaller velocity gradient perpendicular to the filament in the south.

We show in Figure 19 the distributions of σ_v , T_K , $N(\text{NH}_3)$, and $X(\text{NH}_3)$ over the DR1 regions. For each parameter we show the probability density function normalized such that the integral over the range is one for each region. We include only those pixels with small uncertainties in the fitted values, as described in the Figure captions.

Figure 19 (a) reveals a systematic increase in the spread of the NH_3 velocity dispersion in the DR1 regions that reflects the increase in the level of star-formation activity from the least active region, B18, to the most active region, Orion A. In B18, NGC 1333, and L1688, the σ_v distribution shows a narrow peak at small velocity dispersions, and a ‘tail’ of greater velocity dispersions

that increases in both the distribution amplitude relative to the narrow σ_v peak, as well as the maximum σ_v present in the region. In Orion A, the narrow peak might still exist, but the population of larger σ_v values is substantially greater. The lower limit of measured σ_v in all regions is greater than the velocity resolution of the NH_3 data, and will be explored further in a future paper (Pineda et al., in preparation).

Similarly to the distributions of σ_v , both the mean gas kinetic temperature and the spread of temperatures measured varies systematically between the DR1 regions with increasing star-formation activity. Figure 19 (b) shows that B18 contains the coldest dense gas of the DR1 regions, with a mean $T_K \lesssim 10 \text{ K}$ and only a small variation about that value. In both NGC 1333 and L1688, a moderate fraction of the dense gas is similarly cold ($T_K \sim 10 \text{ K}$), but with increasing spread to higher temperatures. Previous pointed NH_3 observations of dense cores in Perseus found a typical temperature of 11 K for these objects, in agreement with our data (Rosolowsky et al. 2008), whereas in NGC 1333 warm gas is concentrated around known young stellar objects. In L1688, dense gas is more likely to be $\sim 14 - 15 \text{ K}$, as seen previously in smaller NH_3 maps of some of the cores in this region (Friesen et al. 2009). Toward Orion A, only a small fraction of the dense gas is as cold as 10 K, with typical temperatures $T_K \sim 15 - 20 \text{ K}$, up to very large values near the Orion KL region (seen in Figure 18). At these temperatures ($>30 \text{ K}$), we expect significant populations of higher-order NH_3 (J,K) levels,

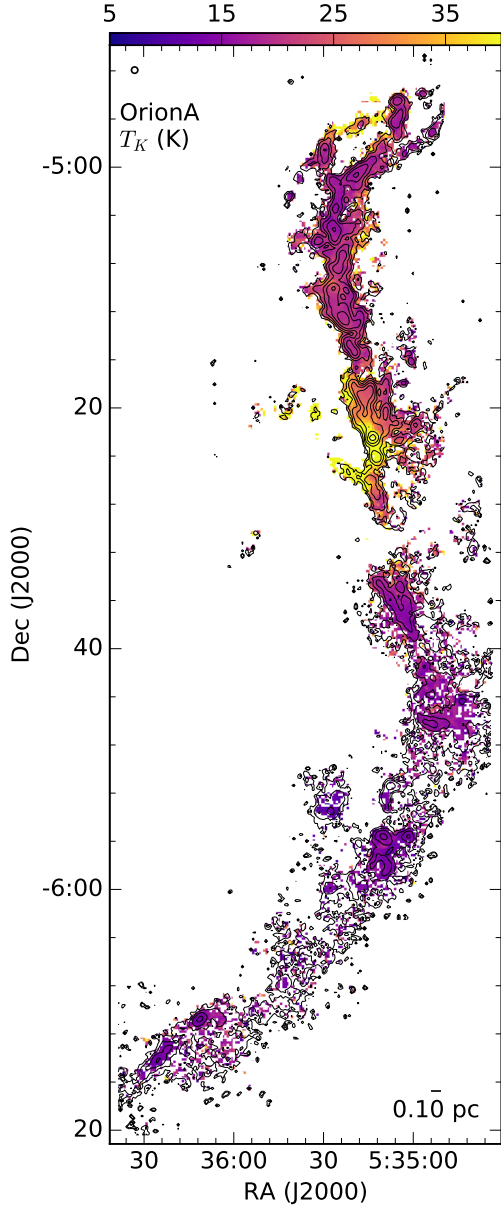


Figure 18. Like Figure 15 but for Orion A (North).

and our temperature estimates from the (1,1) and (2,2) lines become inaccurate.

Figure 19 (c) shows the distribution of NH_3 column densities over the DR1 regions. For all regions, we find a similar increase in $N(p-\text{NH}_3)$ ⁵ at low column densities, and a similar peak at $\log N(p-\text{NH}_3) \sim 14$. B18 shows a relatively sharp cutoff in the maximum column density present. $N(p-\text{NH}_3)$ extends to higher values in the

⁵ In most works the total column density of NH_3 (including ortho- and para-states) is reported, $N(\text{NH}_3)$. A simple way to estimate total column density of NH_3 is: $N(\text{NH}_3) \approx 2 \times N(p-\text{NH}_3)$, if the ortho-to-para ratio is the LTE value of 1.

other regions, however, with the largest column densities found in Orion A.

The resulting distribution of NH_3 abundances, $X(\text{NH}_3) = N(\text{NH}_3)/N(\text{H}_2)$, shown in Figure 19 (d), have very similar peak values of $\log X(\text{NH}_3) \sim -8.5$ to -8.0 . A tail to very high NH_3 abundances is seen in Orion A, but might result from difficulties in accurately fitting the submillimetre continuum SED in this active region. In NGC 1333, the abundance distribution appears bimodal. We will present more detailed analysis of the dust properties and NH_3 abundance distributions in an upcoming paper.

5. SUMMARY

We have presented data from the first release (DR1) for the Green Bank Ammonia Survey, including cubes, moment maps, and property maps, toward B18 in Taurus, L1688 in Ophiuchus, NGC 1333 in Perseus, and Orion A North in the Orion molecular cloud. We furthermore describe in detail the DR1 calibration, imaging, and line fitting pipelines. The four Gould Belt clouds observed span a range of H_2 column density and star formation activity. The extensive and sensitive NH_3 , HC_5N , HC_7N , and C_2S observations of these regions greatly increase the total contiguous areal coverage of dense molecular gas tracers in these well-studied star-forming regions.

All cubes and maps are publicly available through https://dataverse.harvard.edu/dataverse/GAS_DR1, as are the imaging and analysis pipelines.

Some highlights from the sensitive *GAS* data include:

1. Extended NH_3 (1,1) emission is detected toward all DR1 regions, revealing the physical properties of both the dense gas associated with prestellar and star-forming cores and filaments, as well as their moderately dense molecular envelopes.
2. NH_3 (1,1) emission is a particularly good tracer of material at higher $N(\text{H}_2)$, tracking well the H_2 column density as traced by dust continuum emission. In three of four DR1 regions, NH_3 is present above our map rms noise levels in $\gtrsim 60\%$ of pixels at $N(\text{H}_2) \gtrsim 6 \times 10^{21} \text{ cm}^{-2}$, or $A_V \gtrsim 7$.
3. We observe some variation between the regions in the fraction of pixels with NH_3 as H_2 column density increases. We compare the *GAS* data to results calculated from a synthetic observation of NH_3 emission produced by a 3D hydrodynamic simulation. The simulated distributions agree well with the *GAS* distributions and demonstrate that variation in the underlying gas densities can account for some of the spread between regions. Differences between regions may also result from variations in the local UV radiation field and mean column density, which we will explore in future work.

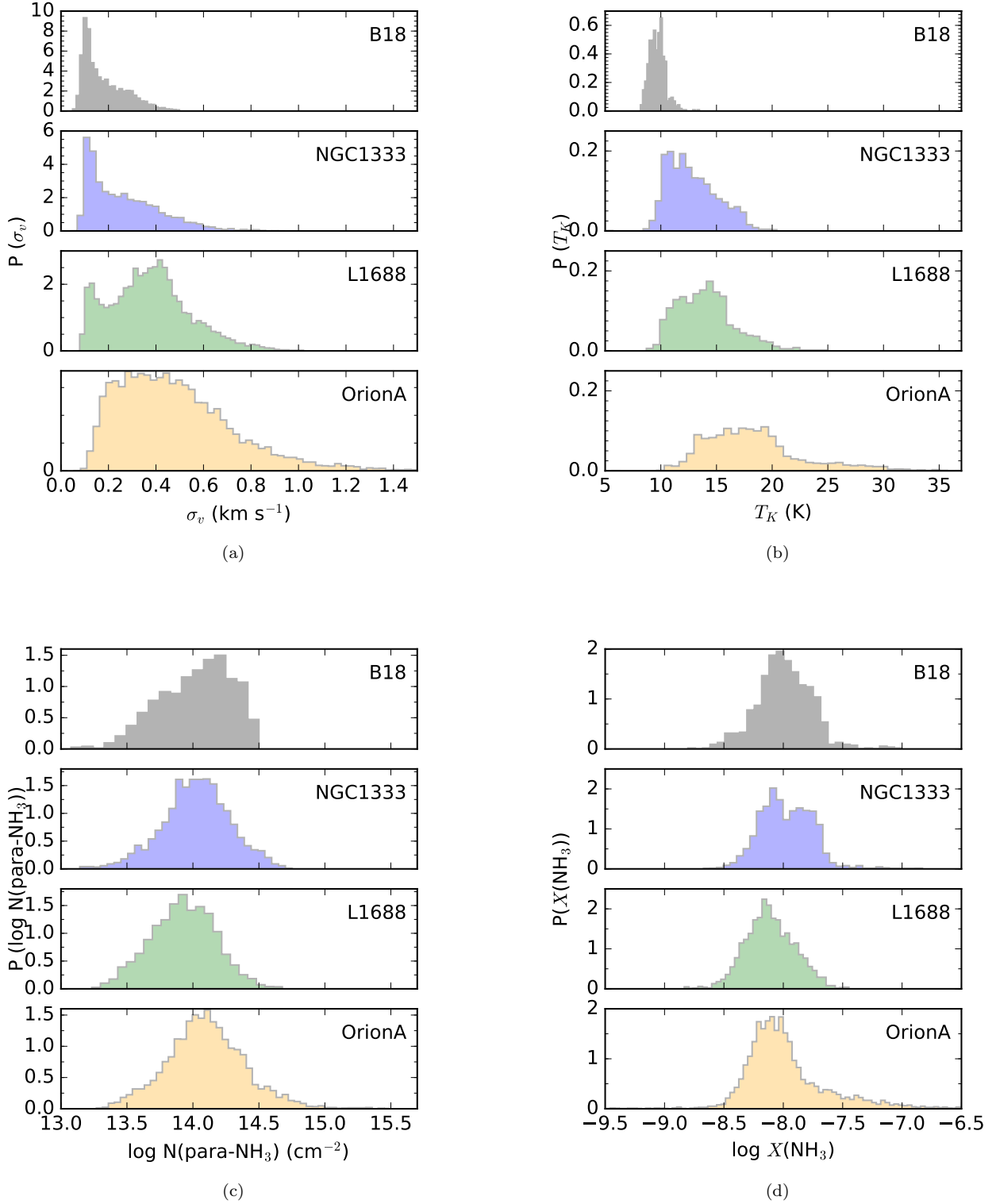


Figure 19. (a) Histograms of σ_v for each DR1 region. Pixels where the NH₃ σ_v is fit with an uncertainty ≤ 0.1 km s⁻¹ are included. In all panels we show the value of the probability density function, normalized such that the integral over the range is 1 for each region. (b) Histograms of T_K for each DR1 region. Pixels where T_K is fit with an uncertainty ≤ 1 K are included. (c) Histograms of $N(\text{NH}_3)$ for each DR1 region. Pixels where $\log N(\text{NH}_3)$ is fit with an uncertainty ≤ 0.25 are included. (d) Histograms of $X(\text{NH}_3)$ for each DR1 region. Pixels where $\log N(\text{NH}_3)$ is fit with an uncertainty ≤ 0.25 are included.

4. Previously unmapped NH_3 (3,3) emission is detected toward NGC 1333, L1688, and Orion A. Highlighting warmer, dense gas, NH_3 (3,3) emission is coincident with YSOs and PDR edges.
5. The carbon-chain molecules C_2S and HC_5N are strongly detected toward Taurus B18, but are faint or undetected in the integrated intensity maps of the other DR1 regions. In Taurus, the spatial distributions of C_2S and HC_5N are offset from the NH_3 emission. We detect C_2S toward NGC 1333 and L1688 only in spectra that were averaged over the entire observation extent. The observed HC_7N lines are not detected toward any DR1 target.
6. The distributions of gas temperature and velocity dispersion from the NH_3 line fits vary systematically between DR1 regions, in step with increasing star formation activity. Lower T_K is found overall in B18, increasing through NGC 1333 and L1688, and is greatest in Orion A. Similarly, σ_v reaches greater values in L1688 and Orion A relative to B18 and NGC 1333.

RKF is a Dunlap Fellow at the Dunlap Institute for Astronomy & Astrophysics. The Dunlap Institute is funded through an endowment established by the David Dunlap family and the University of Toronto. JEP, AP, ACT, and PC acknowledge the financial support of the European Research Council (ERC; project PALs 320620). EWR, PGM and CDM are supported by Discovery Grants from NSERC of Canada. SSRO acknowledges support from NSF grant AST-1510021. The National Radio Astronomy Observatory is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc.

Facility: Green Bank Telescope

Software: Astropy (Robitaille et al. 2013), Matplotlib (Hunter 2007), pyspeckit (Ginsburg & Mirocha 2011), GBTIDL (<http://gbtidl.nrao.edu>), GBT KPFA data reduction pipeline (Masters et al. 2011), RADMC-3D (<http://www.ita.uni-heidelberg.de/~dullemond/software/radmc-3d>)

REFERENCES

- André, P., Francesco, J. D., Ward-Thompson, D., et al. 2014, in *Protostars and Planets VI* (University of Arizona). http://dx.doi.org/10.2458/azu_uapress_9780816531240-ch002
- André, P., Men'shchikov, A., Bontemps, S., et al. 2010, *A&A*, 518, L102
- André, P., Men'shchikov, A., Bontemps, S., et al. 2010, *A&A*, 518, L102. <http://dx.doi.org/10.1051/0004-6361/201014666>
- Arzoumanian, D., André, P., Didelon, P., et al. 2011, *A&A*, 529, L6. <http://dx.doi.org/10.1051/0004-6361/201116596>
- Bally, J. 2008, *Overview of the Orion Complex*, ed. B. Reipurth, 459
- Bally, J., Stark, A. A., Wilson, R. W., & Langer, W. D. 1987, *ApJ*, 312, L45. <http://dx.doi.org/10.1086/184817>
- Benson, P. J., & Myers, P. C. 1983, *ApJ*, 270, 589. <http://dx.doi.org/10.1086/161151>
- Bisbas, T. G., Bell, T. A., Viti, S., Yates, J., & Barlow, M. J. 2012, *MNRAS*, 427, 2100
- Caselli, P., Walmsley, C. M., Tafalla, M., Dore, L., & Myers, P. C. 1999, *ApJ*, 523, L165. <http://dx.doi.org/10.1086/312280>
- Christie, H., Viti, S., Yates, J., et al. 2012, *MNRAS*, 422, 968
- Dobashi, K., Uehara, H., Kandori, R., et al. 2005, *PASJ*, 57, S1
- Dunham, M. M., Stutz, A. M., Allen, L. E., et al. 2014, in *Protostars and Planets VI* (University of Arizona). http://dx.doi.org/10.2458/azu_uapress_9780816531240-ch009
- Dunham, M. M., Allen, L. E., Evans, II, N. J., et al. 2015, *ApJS*, 220, 11
- Enoch, M. L., Evans, N. J., Sargent, A. I., et al. 2008, *ApJ*, 684, 1240. <http://dx.doi.org/10.1086/589963>
- Enoch, M. L., Glenn, J., Evans, II, N. J., et al. 2007, *ApJ*, 666, 982
- Estalella, R. 2017, *PASP*, 129, 025003
- Evans, N. J., Allen, L. E., Blake, G. A., et al. 2003, *PASP*, 115, 965. <http://dx.doi.org/10.1086/376697>
- Evans, N. J., Dunham, M. M., Jørgensen, J. K., et al. 2009, *ApJSS*, 181, 321. <http://dx.doi.org/10.1088/0067-0049/181/2/321>
- Faure, A., Hily-Blant, P., Gal, R. L., Rist, C., & des Forêts, G. P. 2013, *ApJ*, 770, L2. <http://dx.doi.org/10.1088/2041-8205/770/1/L2>
- Foster, J. B., Rosolowsky, E. W., Kauffmann, J., et al. 2009, *ApJ*, 696, 298
- Friesen, R. K., Bourke, T. L., Di Francesco, J., Gutermuth, R., & Myers, P. C. 2016, *ApJ*, 833, 204
- Friesen, R. K., Francesco, J. D., Shirley, Y. L., & Myers, P. C. 2009, *ApJ*, 697, 1457. <http://dx.doi.org/10.1088/0004-637x/697/2/1457>

- Friesen, R. K., Medeiros, L., Schnee, S., et al. 2013, MNRAS, 436, 1513.
<http://dx.doi.org/10.1093/mnras/stt1671>
- Ginsburg, A., & Mirocha, J. 2011, PySpecKit: Python Spectroscopic Toolkit, Astrophysics Source Code Library, , ascl:1109.001
- Goodman, A. A., Barranco, J. A., Wilner, D. J., & Heyer, M. H. 1998, ApJ, 504, 223.
<http://dx.doi.org/10.1086/306045>
- Goodman, A. A., Benson, P. J., Fuller, G. A., & Myers, P. C. 1993, ApJ, 406, 528.
<http://dx.doi.org/10.1086/172465>
- Graves, S. F., Richer, J. S., Buckle, J. V., et al. 2010, MNRAS, 409, 1412. <http://dx.doi.org/10.1111/j.1365-2966.2010.17140.x>
- Griffin, M. J., Abergel, A., Abreu, A., et al. 2010, A&A, 518, L3
- Gutermuth, R. A., Myers, P. C., Megeath, S. T., et al. 2008, ApJ, 674, 336. <http://dx.doi.org/10.1086/524722>
- Hacar, A., Kainulainen, J., Tafalla, M., Beuther, H., & Alves, J. 2016, A&A, 587, A97.
<http://dx.doi.org/10.1051/0004-6361/201526015>
- Hacar, A., & Tafalla, M. 2011, A&A, 533, A34.
<http://dx.doi.org/10.1051/0004-6361/201117039>
- Hacar, A., Tafalla, M., Kauffmann, J., & Kovács, A. 2013, A&A, 554, A55.
<http://dx.doi.org/10.1051/0004-6361/201220090>
- Harvey, P. M., Huard, T. L., Jørgensen, J. K., et al. 2008, The Astrophysical Journal, 680, 495.
<http://dx.doi.org/10.1086/587687>
- Heiderman, A., Evans, II, N. J., Allen, L. E., Huard, T., & Heyer, M. 2010, ApJ, 723, 1019
- Henshaw, J. D., Caselli, P., Fontani, F., et al. 2016a, MNRAS, 463, 146.
<http://dx.doi.org/10.1093/mnras/stw1794>
- Henshaw, J. D., Longmore, S. N., Kruijssen, J. M. D., et al. 2016b, Mon. Not. R. Astron. Soc., 457, 2675.
<http://dx.doi.org/10.1093/mnras/stw121>
- Heyer, M. H. 1988, ApJ, 324, 311.
<http://dx.doi.org/10.1086/165896>
- Heyer, M. H., Vrba, F. J., Snell, R. L., et al. 1987, ApJ, 321, 855. <http://dx.doi.org/10.1086/165678>
- Hildebrand, R. H. 1983, QJRAS, 24, 267
- Hocuk, S., Cazaux, S., Spaans, M., & Caselli, P. 2016, MNRAS, 456, 2586
- Hunter, J. D. 2007, CSE, 9, 90.
<http://dx.doi.org/10.1109/MCSE.2007.55>
- Inutsuka, S., & Miyama, S. M. 1997, ApJ, 480, 681.
<http://dx.doi.org/10.1086/303982>
- Jijina, J., Myers, P. C., & Adams, F. C. 1999, ApJSS, 125, 161. <http://dx.doi.org/10.1086/313268>
- Johnstone, D., Francesco, J. D., & Kirk, H. 2004, ApJ, 611, L45. <http://dx.doi.org/10.1086/423737>
- Jørgensen, J. K., Johnstone, D., Kirk, H., et al. 2008, ApJ, 683, 822. <http://dx.doi.org/10.1086/589956>
- Kirk, H., Johnstone, D., & Basu, S. 2009, ApJ, 699, 1433.
<http://dx.doi.org/10.1088/0004-637X/699/2/1433>
- Kirk, H., Johnstone, D., & Francesco, J. D. 2006, ApJ, 646, 1009. <http://dx.doi.org/10.1086/503193>
- Kirk, H., Myers, P. C., Bourke, T. L., et al. 2013, ApJ, 766, 115
- Koch, E. W., & Rosolowsky, E. W. 2015, MNRAS, 452, 3435
- Könyves, V., André, P., Men'shchikov, A., et al. 2010, A&A, 518, L106.
<http://dx.doi.org/10.1051/0004-6361/201014689>
- Lada, C. J., Lombardi, M., & Alves, J. F. 2010, ApJ, 724, 687
- Li, D., Kauffmann, J., Zhang, Q., & Chen, W. 2013, ApJ, 768, L5.
<http://dx.doi.org/10.1088/2041-8205/768/1/L5>
- Liseau, R., White, G. J., Larsson, B., et al. 1999, A&A, 344, 342
- Loren, R. B. 1989, ApJ, 338, 902.
<http://dx.doi.org/10.1086/167244>
- Loren, R. B., & Wootten, A. 1986, ApJ, 306, 142
- Mangum, J. G., Emerson, D. T., & Greisen, E. W. 2007, A&A, 474, 679.
<http://dx.doi.org/10.1051/0004-6361:20077811>
- Mangum, J. G., & Shirley, Y. L. 2015, PASP, 127, 266.
<http://dx.doi.org/10.1086/680323>
- Maret, S., Faure, A., Scifoni, E., & Wiesenfeld, L. 2009, MNRAS, 399, 425. <http://dx.doi.org/10.1111/j.1365-2966.2009.15294.x>
- Markwardt, C. B. 2009, in Astronomical Society of the Pacific Conference Series, Vol. 411, Astronomical Data Analysis Software and Systems XVIII, ed. D. A. Bohlender, D. Durand, & P. Dowler, 251
- Masters, J., Garwood, B., Langston, G., & Shelton, A. 2011, in Astronomical Society of the Pacific Conference Series, Vol. 442, Astronomical Data Analysis Software and Systems XX, ed. I. N. Evans, A. Accomazzi, D. J. Mink, & A. H. Rots, 127
- Megeath, S. T., Gutermuth, R., Muzerolle, J., et al. 2012, AJ, 144, 192
- Megeath, S. T., Gutermuth, R., Muzerolle, J., et al. 2015, AJ, 151, 5.
<http://dx.doi.org/10.3847/0004-6256/151/1/5>

- Menten, K. M., Reid, M. J., Forbrich, J., & Brunthaler, A. 2007, *A&A*, 474, 515
- Mizuno, A., Onishi, T., Yonekura, Y., et al. 1995, *ApJ*, 445, L161. <http://dx.doi.org/10.1086/187914>
- Myers, P. C., Ladd, E. F., & Fuller, G. A. 1991, *ApJ*, 372, L95. <http://dx.doi.org/10.1086/186032>
- Offner, S. S. R., Bisbas, T. G., Bell, T. A., & Viti, S. 2014, *MNRAS: Letters*, 440, L81. <http://dx.doi.org/10.1093/mnrasl/slu013>
- Onishi, T., Mizuno, A., Kawamura, A., Ogawa, H., & Fukui, Y. 1998, *ApJ*, 502, 296. <http://dx.doi.org/10.1086/305867>
- Ortiz-León, G. N., Loinard, L., Kounkel, M. A., et al. 2017, *ApJ*, 834, 141
- Panopoulou, G. V., Psaradaki, I., Skolidis, R., Tassis, K., & Andrews, J. J. 2017, *MNRAS*, 466, 2529
- Pickett, H. M., Poynter, R. L., Cohen, E. A., et al. 1998, *JQSRT*, 60, 883
- Pineda, J. E., Goodman, A. A., Arce, H. G., et al. 2010, *ApJ*, 712, L116. <http://dx.doi.org/10.1088/2041-8205/712/1/L116>
- . 2011, *ApJ*, 739, L2. <http://dx.doi.org/10.1088/2041-8205/739/1/L2>
- Pineda, J. E., Offner, S. S. R., Parker, R. J., et al. 2015, *Nature*, 518, 213. <http://dx.doi.org/10.1038/nature14166>
- Planck Collaboration XI. 2014, *A&A*, 571, A11
- Press, W. H., Teukolsky, S. A., Vetterling, W. T., & Flannery, B. P. 2002, *Numerical recipes in C++ : the art of scientific computing*
- Rebull, L. M., Padgett, D. L., McCabe, C.-E., et al. 2010, *ApJS*, 186, 259
- Ridge, N. A., Di Francesco, J., Kirk, H., et al. 2006, *AJ*, 131, 2921
- Robitaille, T. P., Tollerud, E. J., Greenfield, P., et al. 2013, *A&A*, 558, A33. <http://dx.doi.org/10.1051/0004-6361/201322068>
- Rosolowsky, E. W., Pineda, J. E., Foster, J. B., et al. 2008, *ApJSS*, 175, 509. <http://dx.doi.org/10.1086/524299>
- Schlaflly, E. F., Green, G., Finkbeiner, D. P., et al. 2014, *ApJ*, 786, 29
- Schneider, S., & Elmegreen, B. G. 1979, *ApJS*, 41, 87. <http://dx.doi.org/10.1086/190609>
- Seo, Y. M., Shirley, Y. L., Goldsmith, P., et al. 2015, *ApJ*, 805, 185. <http://dx.doi.org/10.1088/0004-637X/805/2/185>
- Shirley, Y. L. 2015, *PASP*, 127, 299. <http://dx.doi.org/10.1086/680342>
- Sousbie, T. 2011, *MNRAS*, 414, 350
- Storm, S., Mundy, L. G., Fernández-López, M., et al. 2014, *ApJ*, 794, 165. <http://dx.doi.org/10.1088/0004-637X/794/2/165>
- Stutzki, J., Olberg, M., Winnewisser, G., Jackson, J. M., & Barrett, A. H. 1984, *A&A*, 139, 258
- Swift, J. J., Welch, W. J., & Francesco, J. D. 2005, *ApJ*, 620, 823. <http://dx.doi.org/10.1086/427257>
- Tafalla, M., Myers, P. C., Caselli, P., & Walmsley, C. M. 2004, *A&A*, 416, 191. <http://dx.doi.org/10.1051/0004-6361:20031704>
- Tafalla, M., Myers, P. C., Caselli, P., Walmsley, C. M., & Comito, C. 2002, *ApJ*, 569, 815. <http://dx.doi.org/10.1086/339321>
- van der Tak, F. F. S., Black, J. H., Schöier, F. L., Jansen, D. J., & van Dishoeck, E. F. 2007, *A&A*, 468, 627
- Walsh, A. J., Bourke, T. L., & Myers, P. C. 2006, *ApJ*, 637, 860. <http://dx.doi.org/10.1086/498564>
- Ward-Thompson, D., Di Francesco, J., Hatchell, J., et al. 2007, *PASP*, 119, 855
- Warin, S., Castets, A., Langer, W. D., Wilson, R. W., & Pagani, L. 1996, *A&A*, 306, 935
- Wilking, B. A., Gagné, M., & Allen, L. E. 2008, *Star Formation in the ρ Ophiuchi Molecular Cloud*, ed. B. Reipurth, 351
- Wilking, B. A., & Lada, C. J. 1983, *ApJ*, 274, 698. <http://dx.doi.org/10.1086/161482>

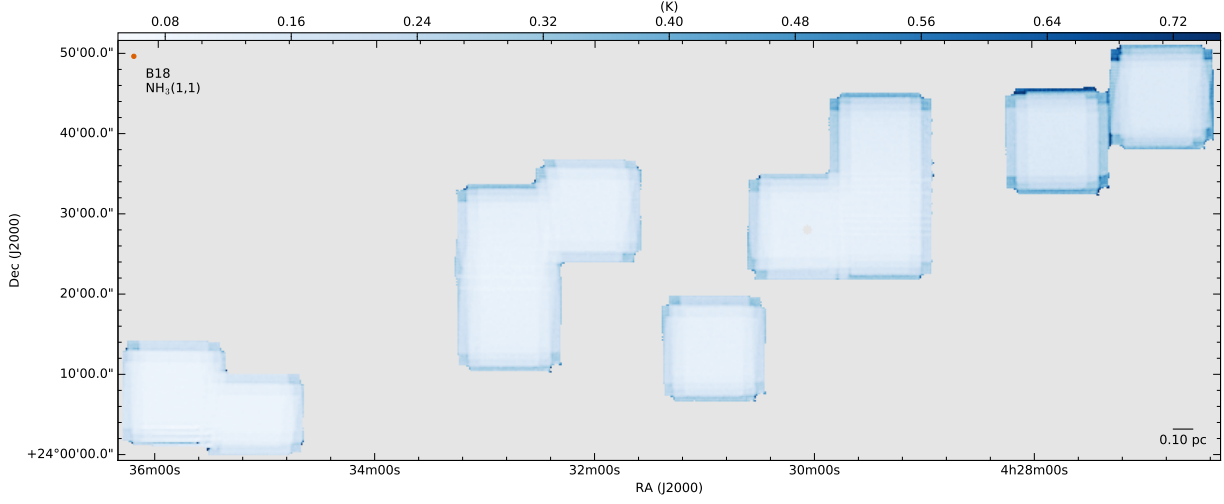


Figure A. The rms noise (K) over the NH_3 (1,1) map toward B18. Scale bar is at lower right.

APPENDIX

A. MAPS OF THE RMS NOISE IN THE NH_3 (1,1) OBSERVATIONS FOR ALL DR1 REGIONS

Fig. Set A. NH_3 (1,1) rms maps

B. NH_3 (2,2) AND (3,3) INTEGRATED INTENSITY MAPS, WHERE DETECTED

Fig. Set B. NH_3 (2,2) Integrated Intensity maps

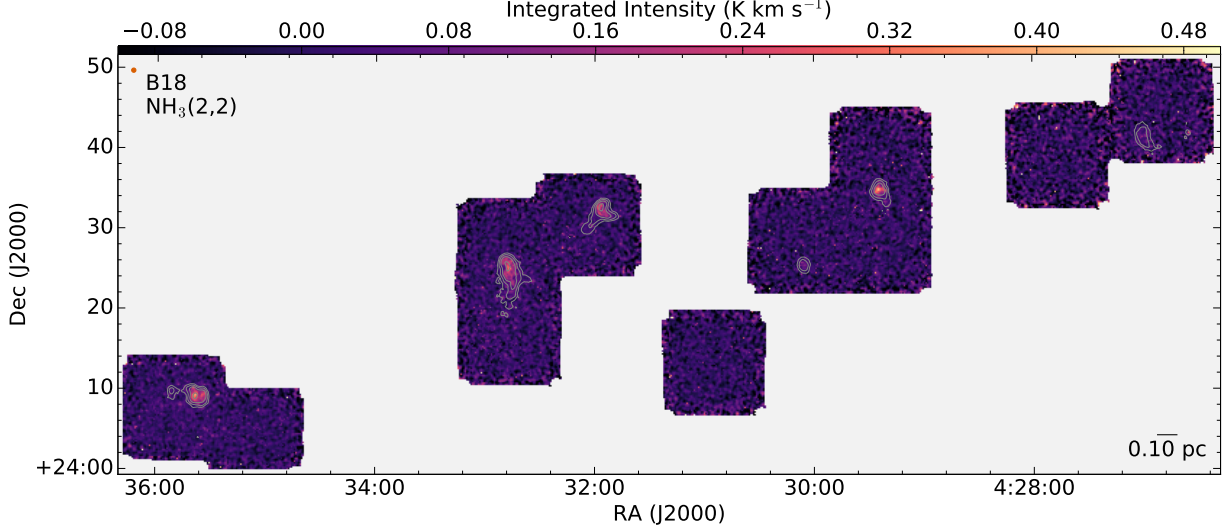


Figure B. The integrated intensity of NH_3 (2,2) toward B18. Contours are NH_3 (1,1) integrated intensity as in Figure 4. Beam size and scale bar are at upper left and lower right, respectively.

Fig. Set C. NH_3 (3,3) Integrated Intensity rms maps for L1688, and Orion A

C. HC_5N 9 – 8 AND C_2S 2₁ – 1₀ INTEGRATED INTENSITY MAPS, WHERE DETECTED

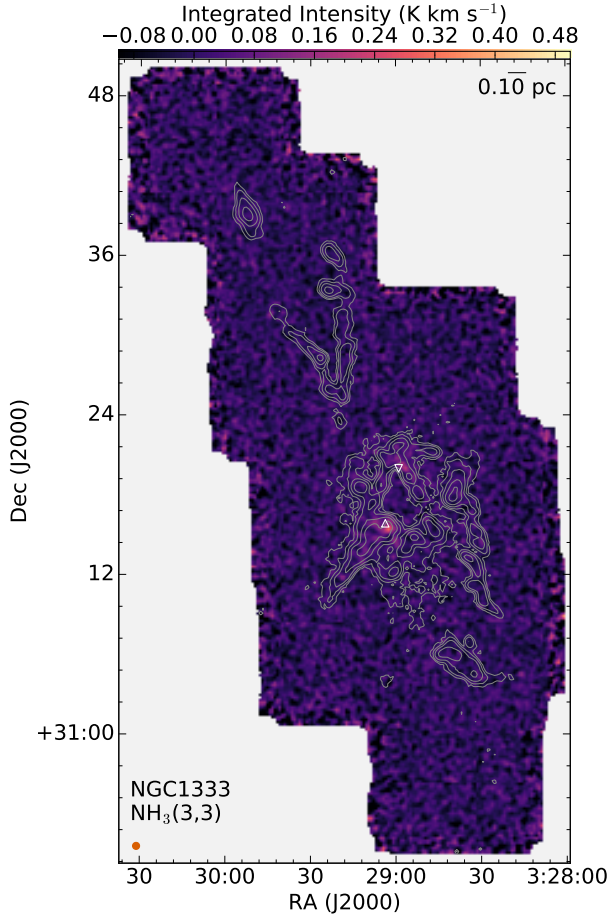


Figure C. NH_3 (3,3) integrated intensity toward NGC 1333. Contours are NH_3 (1,1) integrated intensity as in Figure 5. The locations of the Class 0/I object SVS 13 and Herbig-Haro object HH12 are identified by the upward and downward facing triangles, respectively.

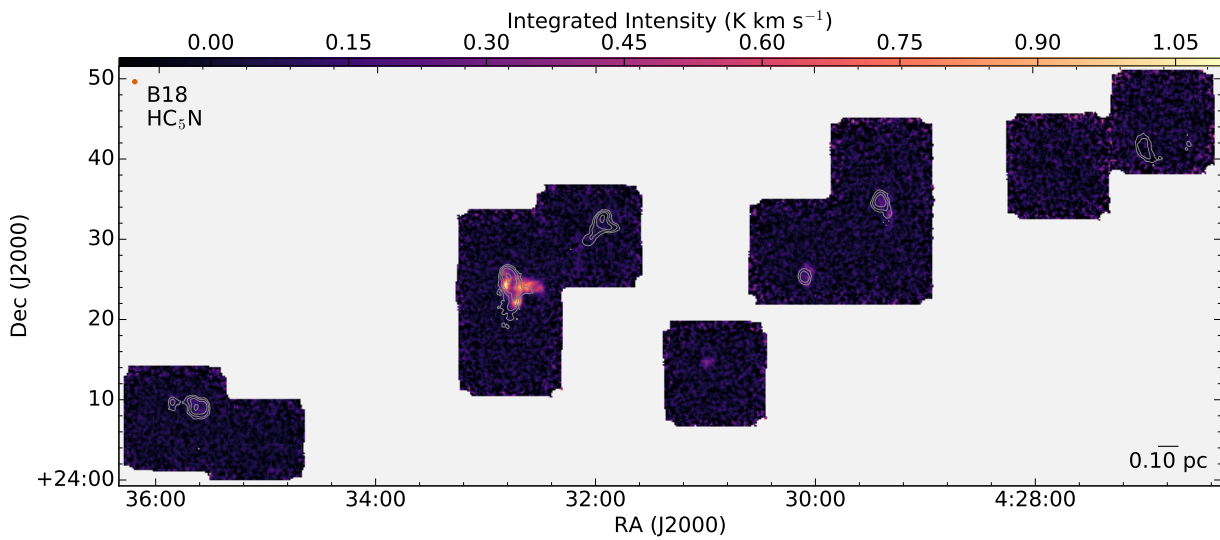


Figure 20. HC_5N 9 – 8 integrated intensity in B18. Contours are NH_3 (1,1) integrated intensity as in Figure 4.

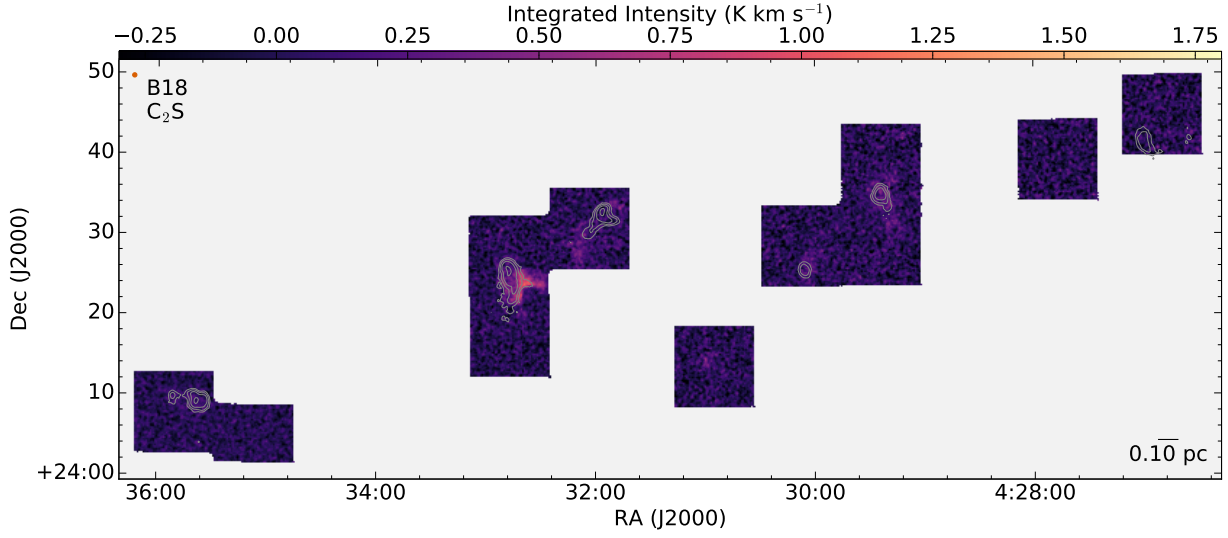


Figure 21. C_2S $2_1 - 1_0$ integrated intensity toward B18. Contours are NH_3 (1,1) integrated intensity as in Figure 4. Note that C_2S was observed in a single beam only, and hence has a smaller footprint and higher noise values than the other mapped lines, but remains Nyquist-sampled.

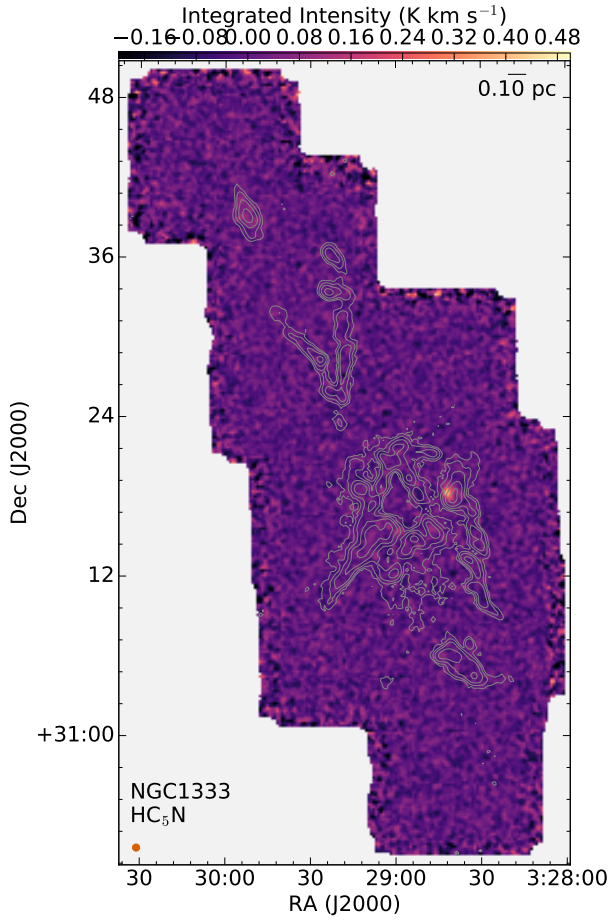


Figure 22. Like Figure 20 but for NGC 1333.

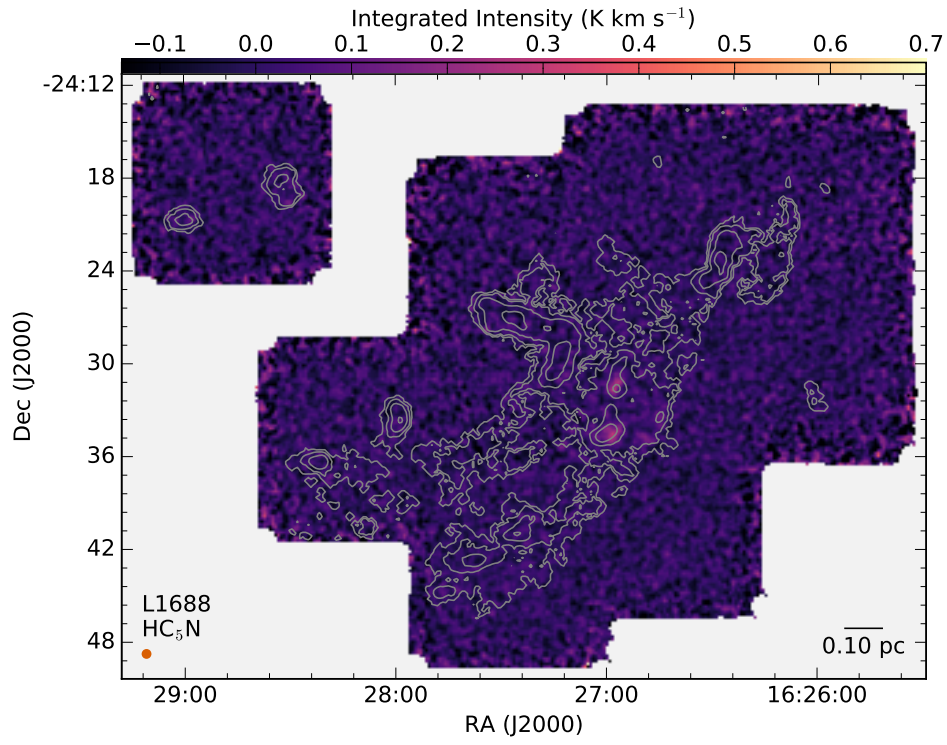


Figure 23. Like Figure 20 but for L1688.

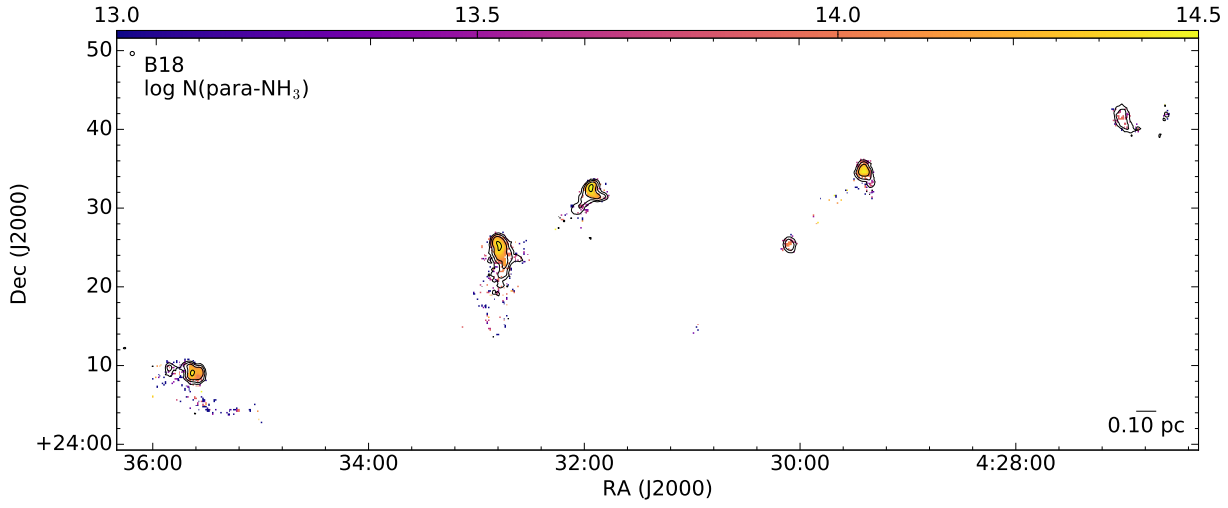
D. NH_3 COLUMN DENSITY MAPS FOR ALL DR1 REGIONS

Figure 24. $N(\text{NH}_3)$ (cm^{-2}) in B18. Contours are NH_3 (1,1) integrated intensity as in Figure 4.

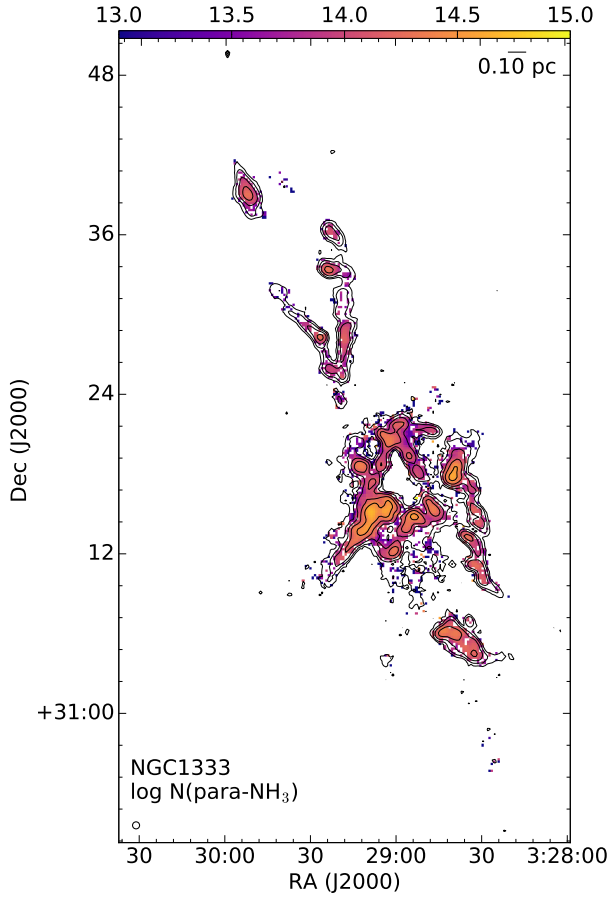


Figure 25. Like Figure 24 but for NGC 1333.

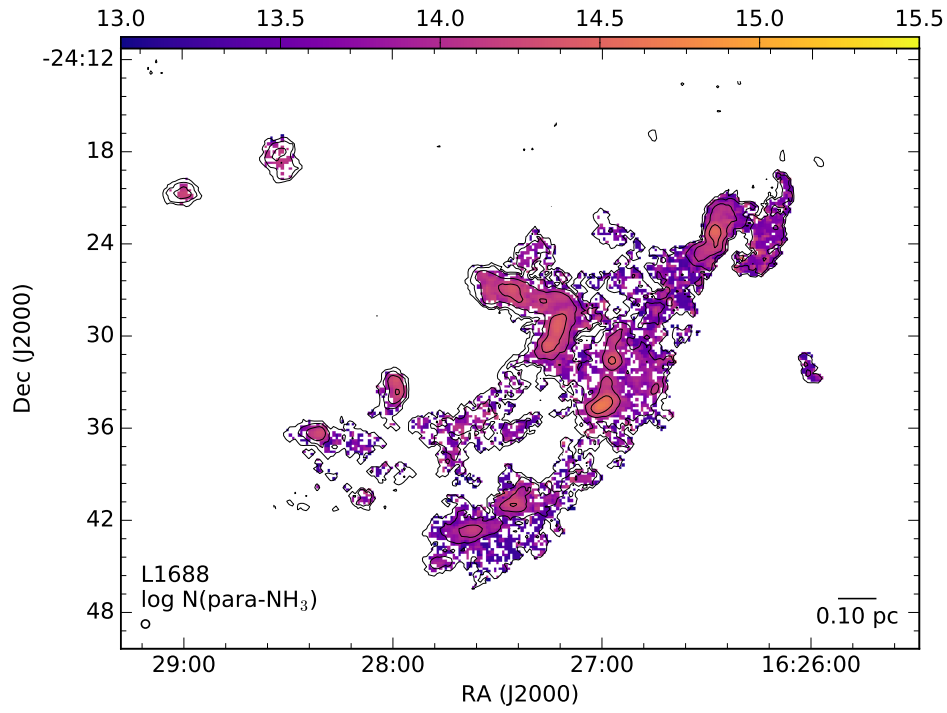


Figure 26. Like Figure 24 but for L1688.

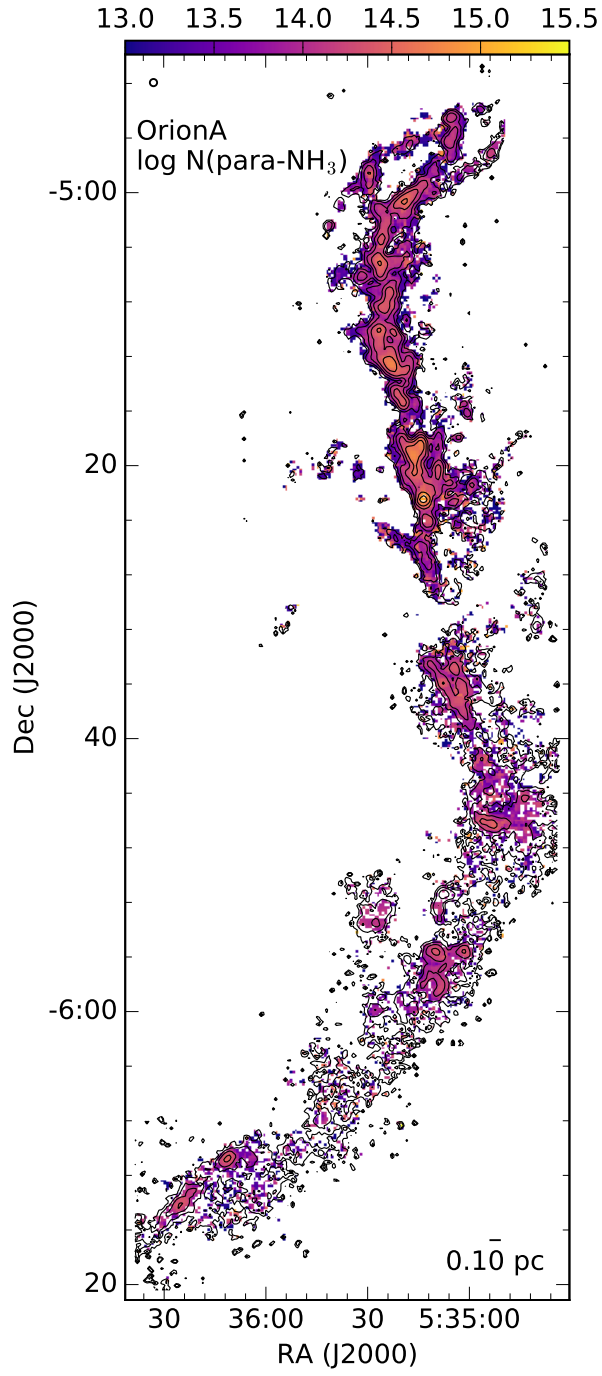


Figure 27. Like Figure 24 but for Orion A (North).